



THE MUSCLES TREASURY SURVEY. III. X-RAY TO INFRARED SPECTRA OF 11 M AND K STARS HOSTING PLANETS

R. O. P. LOYD¹, KEVIN FRANCE¹, ALLISON YOUNGBLOOD¹, CHRISTIAN SCHNEIDER², ALEXANDER BROWN³, RENYU HU^{4,5}, JEFFREY LINSKY⁶, CYNTHIA S. FRONING⁷, SETH REDFIELD⁸, SARAH RUGHEIMER⁹, AND FENG TIAN¹⁰

¹Laboratory for Atmospheric and Space Physics, Boulder, Colorado, 80309, USA; robert.loyd@colorado.edu

²European Space Research and Technology Centre (ESA/ESTEC), Keplerlaan 1, 2201 AZ Noordwijk, The Netherlands

³Center for Astrophysics and Space Astronomy, University of Colorado, 389 UCB, Boulder, CO 80309, USA

⁴Jet Propulsion Laboratory, California Institute of Technology, Pasadena, CA 91109, USA

⁵Division of Geological and Planetary Sciences, California Institute of Technology, Pasadena, CA 91125, USA

⁶JILA, University of Colorado and NIST, 440 UCB, Boulder, CO 80309, USA

⁷Dept. of Astronomy C1400, University of Texas, Austin, TX, 78712, USA

⁸Astronomy Department and Van Vleck Observatory, Wesleyan University, Middletown, CT 06459-0123, USA

⁹Dept of Earth and Environmental Sciences, University of St. Andrews, Irvine Building, North Street, St. Andrews KY16 9AL, UK

¹⁰Ministry of Education Key Laboratory for Earth System Modeling, Center for Earth System Science, Tsinghua University, Beijing, 100084, China

Received 2016 January 28; accepted 2016 April 12; published 2016 June 20

ABSTRACT

We present a catalog of panchromatic spectral energy distributions (SEDs) for 7 M and 4 K dwarf stars that span X-ray to infrared wavelengths (5 Å – 5.5 μm). These SEDs are composites of *Chandra* or *XMM-Newton* data from 5–~50 Å, a plasma emission model from ~50–100 Å, broadband empirical estimates from 100–1170 Å, *Hubble Space Telescope* data from 1170–5700 Å, including a reconstruction of stellar Ly α emission at 1215.67 Å, and a PHOENIX model spectrum from 5700–55000 Å. Using these SEDs, we computed the photodissociation rates of several molecules prevalent in planetary atmospheres when exposed to each star’s unattenuated flux (“unshielded” photodissociation rates) and found that rates differ among stars by over an order of magnitude for most molecules. In general, the same spectral regions drive unshielded photodissociations both for the minimally and maximally FUV active stars. However, for O₃ visible flux drives dissociation for the M stars whereas near-UV flux drives dissociation for the K stars. We also searched for an far-UV continuum in the assembled SEDs and detected it in 5/11 stars, where it contributes around 10% of the flux in the range spanned by the continuum bands. An ultraviolet continuum shape is resolved for the star ϵ Eri that shows an edge likely attributable to Si II recombination. The 11 SEDs presented in this paper, available online through the Mikulski Archive for Space Telescopes, will be valuable for vetting stellar upper-atmosphere emission models and simulating photochemistry in exoplanet atmospheres.

Key words: stars: low-mass – ultraviolet: stars – X-rays: stars

1. INTRODUCTION

Current stellar and planetary population statistics indicate that most rocky planets orbit low-mass stars. Low-mass stars, specifically spectral types M and K, greatly outnumber those of higher mass, making up at least 91% of the stellar population within 10 pc of the Sun (Henry et al. 2006). On average, the low-mass stellar population exhibits a planetary occurrence rate of 0.1–0.6 habitable zone terrestrial planets per star (Dressing & Charbonneau 2013; Kopparapu 2013; Dressing & Charbonneau 2015). Further, the occurrence rate of rocky planets was found to decrease with increasing stellar mass by Howard et al. (2012) and no trend with stellar mass was found by Fressin et al. (2013). The above results collectively imply that, by numbers alone, low-mass stars are certain to be a cornerstone of exoplanet science and the search for other Earths.

The abundance of low-mass stars ensures that many are close enough to enable high-precision photometry and spectroscopy. In addition, their smaller sizes and lower masses yield deeper planetary transits and larger stellar reflex radial velocities when compared to a system with a higher stellar mass but identical planetary mass and orbital period. Furthermore, orbits around cooler, less-luminous stars must have shorter periods than those around Sun-like stars to achieve the same planetary effective temperature, making transits more likely and frequent and enhancing reflex velocities for radial velocity detection. These

advantages facilitate the detection and bulk characterization (e.g., mass, radius) of planets orbiting nearby low-mass stars using radial velocity and transit techniques. They also facilitate atmospheric characterization through transmission spectroscopy, as has recently been performed for super-Earths orbiting the M4.5 star GJ 1214 and K1 star HD 97658 (Knutson et al. 2014; Kreidberg et al. 2014).

Upcoming exoplanet searches, such as the *Transiting Exoplanet Survey Satellite* (*TESS*), will focus on low-mass stars. In searching for planets as cool as Earth, *TESS* is biased by its short (1–12 month) monitoring of host stars (Ricker et al. 2014). Only around low-mass stars will cool planets orbit with short enough periods to transit several times during these monitoring programs. As a result, the *TESS* sample of potentially habitable exoplanets will mostly be orbiting low-mass stars (Deming et al. 2009).

The prevalence of low-mass stars hosting planets makes a thorough knowledge of the typical planetary environment provided by such stars indispensable. However, the circumstellar environment of an M or a K dwarf differs substantially from the well-studied environment of the Sun. Lower-mass stars have cooler photospheres that emit spectra peaking at redder wavelengths in comparison with Sun-like stars, but there are other important differences to address. As we discuss these differences and throughout this paper, we will refer to various

ranges of the stellar spectral energy distributions (SEDs) using their established monikers. We adopt the definitions of X-ray $<100 \text{ \AA}$, EUV = $[100 \text{ \AA}, 912 \text{ \AA}]$ (912 \AA corresponds to the red edge of emission from H II recombination), XUV (X-ray + EUV) $< 912 \text{ \AA}$, far-UV (FUV) = $[912 \text{ \AA}, 1700 \text{ \AA}]$, near-UV (NUV) = $[1700 \text{ \AA}, 3200 \text{ \AA}]$, visible = $[3200 \text{ \AA}, 7000 \text{ \AA}]$, IR $> 7000 \text{ \AA}$. Although emission from the stellar photosphere dominates the energy budget of emitted stellar radiation ($F(<1700 \text{ \AA})/F(>1700 \text{ \AA}) \approx 10^{-4}$ for the stars in this paper), X-ray through UV emission that is emitted primarily from the stellar chromosphere and corona drives photochemistry, ionization, and mass loss in the atmospheres of orbiting planets. We briefly describe these influences below.

FUV photons are able to dissociate molecules, both heating the atmosphere and directly modifying its composition. This affects many molecules common in planetary atmospheres, including O_2 , H_2O , CO_2 , and CH_4 . If the fraction of FUV to bolometric flux is sufficiently large, photodissociation can push the atmosphere significantly out of thermochemical equilibrium (see, e.g., Hu & Seager 2014; Miguel & Kaltenegger 2014; Moses 2014). Photochemical reactions can produce a buildup of O_2 and O_3 at concentrations that, in the absence of radiative forcing, might be considered indicators of biological activity (Domagal-Goldman et al. 2014; Tian et al. 2014; Wordsworth & Pierrehumbert 2014; Harman et al. 2015). It is also possible that these atmospheric effects, besides just interfering with the detection (or exclusion) of existing life, might also influence the emergence and evolution of life. UV radiation can both damage (Voet et al. 1963; Matsunaga et al. 1991; Tevini 1993; Kerwin & Remmele 2007) and aid in synthesizing (Senanayake & Idriss 2006; Barks et al. 2010; Ritson & Sutherland 2012; Patel et al. 2015) many molecules critical to the function of Earth’s life.

At higher photon energies, stellar extreme ultraviolet (EUV) and X-ray photons (often termed XUV in conjunction with the EUV) can eject electrons from atoms, ionizing and heating the upper atmospheres of planets. This heating can drive significant atmospheric escape for close-in planets (Lammer et al. 2003; Yelle 2004; Tian et al. 2005; Murray-Clay et al. 2009). In addition, the development of an ionosphere through ionization by EUV flux will influence the interaction (and associated atmospheric escape) of the planetary magnetosphere with the stellar wind (e.g., Cohen et al. 2014). Indeed, atmospheric escape has been observed on the hot Jupiters HD 209458b (Vidal-Madjar et al. 2003; Linsky et al. 2010), HD 189733b (Lecavelier Des Etangs et al. 2010), and WASP-12b (Fossati et al. 2010, 2013) and the hot Neptune GJ 436b (Kulow et al. 2014; Ehrenreich et al. 2015).

The effects of stellar radiation on planetary atmospheres warrant the spectroscopic characterization of low-mass stars at short wavelengths. Yet these observations are rare compared to visible-IR observations and model-synthesized spectra. To address this scarcity, France et al. (2013) conducted a pilot program collecting UV spectra of six M dwarfs, a project termed Measurements of the Ultraviolet Spectral Characteristics of Low-Mass Exoplanetary Systems (MUSCLES).

Responding to the success of the pilot survey and continued urging of the community (e.g., Segura et al. 2005; Domagal-Goldman et al. 2014; Tian et al. 2014; Cowan et al. 2015; Rugheimer et al. 2015), we have completed the MUSCLES Treasury Survey, doubling the stellar sample and expanding the spectral coverage from $\Delta\lambda \approx 2000 \text{ \AA}$ in the ultraviolet to a

span of four orders of magnitude in wavelength ($5 \text{ \AA} - 5.5 \mu\text{m}$). We have combined observations in the X-ray, UV, and blue-visible; reconstructions of the ISM-absorbed Ly α and EUV flux; and spectral output of the PHOENIX photospheric models (Husser et al. 2013; $5700 \text{ \AA} - 5.5 \mu\text{m}$) and APEC coronal models (Smith et al. 2001; $\sim 50 - 100 \text{ \AA}$) to create panchromatic spectra for all targets. The fully reduced and coadded spectral catalog is publicly available through the Mikulski Archive for Space Telescopes (MAST).¹¹

We present the initial results of the MUSCLES Treasury Survey in three parts: France et al. (2016; hereafter Paper I) gives the overview of this MUSCLES Treasury program including an investigation of the dependence of total UV and X-ray luminosities and individual line fluxes on stellar and planetary parameters that reveals a tantalizing suggestion of star–planet interactions. Youngblood et al. (2016; hereafter “Paper II”) supplies the details of the Ly α reconstruction and EUV modeling and explores the possibility of estimating EUV and Ly α fluxes from other, more readily observed emission lines and the correlation of these fluxes with stellar rotation. This paper, the third in the series, discusses the details of assembling the panchromatic SEDs that are the primary data product of the MUSCLES Treasury Survey, with a report on the detection of FUV continua in the targets, a computation of “unshielded” dissociation rates for important molecules in planetary atmospheres, and a comparison to purely photospheric PHOENIX spectra.

This paper is structured as follows: Section 2 describes the data reduction process, separately addressing each of the sources of the spectra that are combined into the composite SEDs, the process for this combination, special cases, and overall data quality. Section 3 explores the SED catalog and discusses its use. There we examine the FUV continuum, make suggestions for estimating the SEDs of non-MUSCLES stars, compute and compare unattenuated photodissociation rates, and illustrate the importance of accounting for emission from stellar upper atmospheres in SEDs. Section 4 then summarizes the data products and results.

2. THE DATA

Assembling panchromatic SEDs for the MUSCLES spectral catalog requires data from many sources, including both observations and models. Observational data was obtained using the space telescopes *Hubble Space Telescope* (HST), *Chandra*, and *XMM-Newton* through dedicated observing programs. Model output was obtained from APEC plasma models (Smith et al. 2001), empirical EUV predictions (Paper II), Ly α reconstructions (Paper II), and PHOENIX atmospheric models (Husser et al. 2013). The approximate wavelength range covered by each source is illustrated in Figure 1. We describe each data source and the associated reduction process in order of increasing wavelength below, followed by a discussion of how these sources were joined to create panchromatic SEDs. Paper I includes further details on the rationale and motivation for the various sources.

Two versions of the SED data product, one that retains the native source resolutions and another rebinned to a constant 1 \AA resolution, are available for download through the MAST High Level Science Product Archive.¹² We made one

¹¹ <https://archive.stsci.edu/prepds/muscles/>

¹² <https://archive.stsci.edu/prepds/muscles/>

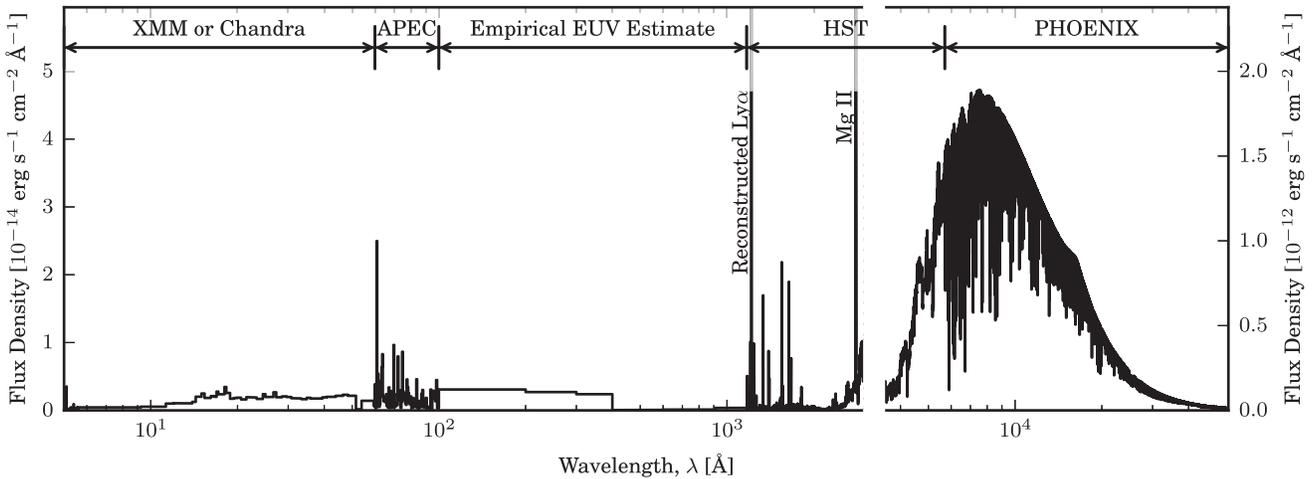


Figure 1. Source data ranges for the MUSCLES composite panchromatic SEDs shown with the GJ 832 spectrum for reference. The gap in the axes indicates a change in the vertical plot scale. Both the Ly α and Mg II lines extend beyond the vertical range of the plot.

exception to retaining native source resolutions by upsampling the broad (100 Å) bands of the EUV estimates to 1 Å binning to ensure accuracy for users that numerically integrate the spectra using the bin midpoints. We note, however, that the most precise integration over any portion of the SED is obtained through multiplying bin widths by bin flux densities and summing.

The detailed format of this data product is described in the readme file included in the archive, permitting small improvements to be made over time and reflected in the living readme document. The spectral data products retain several types of relevant information, beyond just the flux density in each bin. This information is propagated through the data reduction pipeline at each step. As a result, every spectral bin of the final SEDs provides (at the time of publication) information on the

1. Bin: wavelength of the edges and midpoint (Å).
2. Flux Density: measurement and error of the absolute value ($\text{erg s}^{-1} \text{cm}^{-2} \text{Å}^{-1}$) along with the value normalized by the estimated bolometric flux (Å^{-1}) (this estimate is discussed in Section 2.7.4).
3. Exposure: Modified Julian Day of the start of the first contributing exposure, end of the last contributing exposure, and the cumulative exposure time (s).
4. Normalization: any normalization factor applied to the data prior to splicing into the composite SED (applies only to PHOENIX model and HST STIS data, see Sections 2.3 and 2.4).
5. Data source: a bit-wise flag identifying the source of the flux data for the bin.
6. Data quality: bit-wise flags of data quality issues (HST data only).

2.1. X-Ray Data

We obtained X-ray data with *XMM-Newton* for GJ 832, HD 85512, HD 40307, and ϵ Eri and with *Chandra* for GJ 1214, GJ 876, GJ 581, GJ 436, and GJ 176. Complete details of the X-ray data reduction and an analysis of the results will be presented in a follow-on paper (A. Brown et al. 2016, in preparation). Here we provide a brief outline of the reduction process. The technique used to observe X-ray photons imposes

some challenges to extracting the source spectrum from the observational data. The X-ray emission was observed using CCD detectors that have much lower spectral resolution ($\lambda/\Delta\lambda \sim 10\text{--}15$ for the typical 1 keV photons observed from the MUSCLES stars) than the spectrographs used to measure the UV and visible spectra. The data were recorded in instrumental modes designed to allow the efficient rejection of particle and background photon events. CCD X-ray detectors have very complex instrumental characteristics that are energy-dependent and also depend critically on where on the detector a source is observed. Broad, asymmetric photon energy contribution functions result from these effects, and it is nontrivial to associate a measured event energy to the incident photon energy. Consequently, simply assigning to each detection the most likely energy of the photon that created the detected event and binning in energy does not accurately reproduce the source spectrum.

To estimate the true source spectrum, we used the XSPEC software package (Arnaud 1996) to forward model and parameterize the observed event list, using model spectra generated by Astrophysical Plasma Emission Code (APEC; Smith et al. 2001). For each spectrum, we experimented with fits to plasma models containing one to three temperature components and with elemental abundance mixes that were either fixed to solar values or free to vary (but normally only varying the Fe abundance, because it is the dominant emission line contributor) until we achieved the most reasonable statistical fit. The complexity of the model fit is strongly controlled by the number of detected X-ray events. The resulting model parameters are listed in Table 1.

We then computed the ratio of the number of photons incident in each energy bin to the number actually recorded, and corrected the measured photon counts accordingly. The spectral bin widths are variable because the data were binned to contain an equal number of counts per energy bin.

2.2. EUV and Ly α

Much of the EUV radiation emitted by a star, along with the core of the Ly α line, cannot be observed because of absorption and scattering by hydrogen in the interstellar medium. However, emission in the wings of the Ly α line does reach

Table 1
Parameters of the APEC Model Fits to the X-Ray Data

Star	kT^a (keV)	EM_i/EM_1^b	F_x^c (10^{-14} erg s $^{-1}$ cm $^{-2}$)
GJ 1214	0.2 ^d		≤ 0.11
GJ 876	0.80 \pm 0.14 0.14 \pm 0.04	3.5 \pm 4.6	9.1 \pm 0.8
GJ 436	0.39 \pm 0.03		1.2 \pm 0.1
GJ 581	0.26 \pm 0.02		1.8 \pm 0.2
GJ 667C	0.41 \pm 0.03		3.9 \pm 0.3
GJ 176	0.31 \pm 0.02		4.8 \pm 0.3
GJ 832	0.38 ^{+0.11} _{-0.07} 0.09 ^{+0.02} _{-0.09}	4.8 ^{+2.6} _{-2.9}	6.2 ^{+0.8} _{-0.7}
HD 85512	0.25 ^{+0.04} _{-0.03}		1.9 ^{+0.4} _{-0.3}
HD 40307	0.15 \pm 0.06		1.0 \pm 0.2
HD 97658 ^e			
ϵ Eri	0.70 ^{+0.04} _{-0.07} 0.32 \pm 0.01 0.12 ^{+0.09} _{-0.02}	3.6 ^{+0.7} _{-0.9} 2.4 ^{+2.0} _{-1.0}	940 \pm 10

Notes.

^a Plasma temperature. Multiple values for an entry represent multiple plasma components in the model.

^b For multi-component plasmas, this represents the ratio of the emission measure of the i th component to the first component according to the order listed in the kT column.

^c Flux integrated over full instrument bandpass.

^d Estimate is not well constrained and is based on an earlier *XMM-Newton* observation. See Section 2.6.

^e No observations. See Section 2.6.

Earth because multiple scatterings in the stellar atmosphere produce broad emission wings that are not affected by the narrower absorption profile produced by H I and D I in the ISM. This allows the full line to be reconstructed by modeling these processes.

Youngblood et al. (2016; Paper II) have reconstructed the intrinsic Ly α profile from observations made with *HST* STIS G140M and (for bright sources) E140M data. The narrow slits used to collect these data ($52'' \times 0.''1$ for G140M and $0.''2 \times 0.''06$ for E140M) produce a spectrum where the diffuse geocoronal Ly α “airglow” emission is spatially extended beyond the target spectrum and spectrally resolved. This allows the geocoronal spectrum to be measured and subtracted from the observed spectrum, leaving only the target flux. The modeling procedure used to reconstruct the intrinsic Ly α line from the airglow-subtracted data is described in detail in Paper II. This reconstruction covers 1209.5 to 1222.0 Å in the MUSCLES spectra.

The EUV is estimated from the intrinsic Ly α flux for each source using the empirical fits of Linsky et al. (2014). This also is detailed in Paper II. We use these estimates to fill all of the EUV as well as the portion of the FUV below 1170 Å where the reflectivity of the Al+MgF $_2$ coatings in the *HST* spectrograph optics rapidly declines.

2.3. FUV through Blue-visible

Currently, *HST* is the only observatory that can obtain spectra at UV wavelengths. We used *HST* with the COS and STIS spectrographs to obtain UV data of all 11 sources. Obtaining full coverage of the UV required multiple observations using complementary COS and STIS gratings. The choice

of gratings depended on the target brightness. Figure 2 illustrates these configurations, depicting which instrument and which grating provided the data for the different pieces of each star’s UV dataset. Along with the UV observations, we also obtained a visible spectrum with STIS G430L covering visible wavelengths up to 5700 Å (with the exception of ϵ Eri, for which instrument brightness limits required the use of STIS G430M covering ~ 3800 – 4075 Å) since the required observing time once the telescope was already pointed was negligible.

For observations with multiple exposures, we coadded data from each exposure. We did the same for the (overlapping) orders of echelle spectra. This produced a single spectrum for each instrument configuration.

We inspected all the coadded spectra and culled any data suspected of detector edge effects. We also removed emission from the geocoronal airglow present in the COS G130M spectra if the airglow line was visible in the spectrum of at least one MUSCLES source. The wavelengths at which we inspected the MUSCLES data for this emission are those listed in the¹³ compiled from airglow-only observations. The removed lines were N I $\lambda 1134$, He I $\lambda 584$ at second order (1168 Å), N I $\lambda 1200$, O I $\lambda 1305$, and O I $\lambda 1356$. The resulting gaps were later filled with STIS data where there was overlap and a quadratic fit to the nearby continuum where there was not.

2.3.1. A Discrepancy in the Absolute Level of Flux Measurements by COS and STIS

Several instrument configurations of the *HST* data have overlapping wavelength ranges. This allowed us to compare the fluxes from these configurations prior to stitching the spectra into the final panchromatic SED. Where there was sufficient signal for a meaningful comparison, STIS always measured lower fluxes than COS by factors of 1.1–2.4. The cause of this discrepancy could be a systematic inaccuracy in the STIS data, the COS data, or both.

The STIS G430L data can be compared against external data in search of a systematic trend because the grating bandpass overlaps with the standard B band for which many ground-based photometric measurements are available. Carrying out this comparison showed that the fluxes measured with the STIS G430L grating were lower than ground-based B-band photometry for every star. The magnitude of these discrepancies varied from the discrepancies between STIS and COS data, but not beyond what is reasonable given uncertainties in the B-band photometry. A plausible cause of systematically low-flux measurements by STIS could be imperfect alignment of the spectrograph slit on the target. This can produce significant flux losses when a narrow slit is used (Biretta et al. 2015; see Section 13.7.1). No such comparison with external data was possible for the COS data. While there is overlap of COS data with *GALEX* bands, *GALEX* photometry is only available for roughly half the targets and uncertainties are very large.

Given the low STIS fluxes relative to ground-based photometry, the plausible explanation of such low fluxes, and the lack of an external check for the COS absolute flux accuracy, we chose to treat the absolute flux levels to be accurate for COS and inaccurate for STIS. Thus, for each STIS spectrum, we normalized to overlapping COS data whenever the difference in flux in regions of high signal was sufficiently large to admit a $<5\%$ false alarm probability. This condition

¹³ http://www.stsci.edu/hst/cos/calibration/airglow_table.html

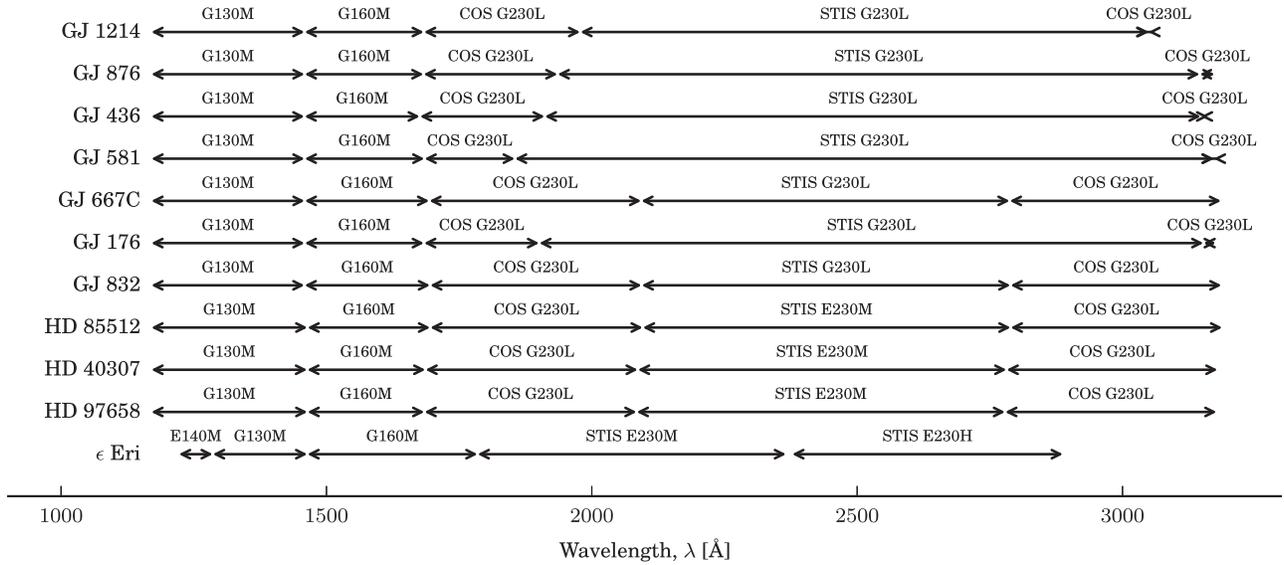


Figure 2. Sources of the UV data for the composite SED of each MUSCLES star. Due to space constraints, the labels E140M, G130M, and G160M are used in place of STIS E140M, COS G130M, and COS G160M. Note that the STIS G140M or E140M data obtained for each star do not appear on this figure because they were only used to reconstruct the Ly α emission (Section 2.2) and fill small gaps left by airglow removal (Section 2.3).

was met for all of the G230L and E230M spectra (excepting ϵ Eri for which comparable COS NUV data could not be collected due to overlap concerns), a few G140M spectra, and the E140M spectrum of ϵ Eri. The normalization factor was simply computed as the ratio of the integrated fluxes in the chosen region. We normalized data from each instrumental configuration separately, rather than normalizing all STIS data by the same factor, to allow for potential variations in throughput between instrument configurations. An example of this normalization is displayed in Figure 3.

The STIS G430L spectra have a small amount of overlap with COS G230L near 3100 Å that could be used for normalization. However, the G430L data are of poor quality in this region, making normalization factors very uncertain and heavily dependent on the wavelength range used. Thus, we normalized the G430L data to all available non-MUSCLES photometry via a PHOENIX model fit, detailed in the following section.

The normalization factors applied are recorded pixel-by-pixel as a separate column in the final data files.

2.4. Visible through IR

We used synthetic spectra generated by Husser et al. (2013) from a PHOENIX stellar atmosphere model to fill the range from the blue-visible at ~ 5700 Å out to $5.5 \mu\text{m}$ where their model spectra truncate. The choice to use model output instead of observations enabled greater consistency in the treatment of the visible and IR between sources, given that for some sources we were unable to acquire optical and IR spectra within the same time window as the X-ray and UV observations. The visible and IR emission of low-mass stars is well reproduced by PHOENIX models. The Husser et al. (2013) PHOENIX spectra cover a grid in effective temperature (T_{eff}), surface gravity ($\log_{10} g$), metallicity ($[\text{Fe}/\text{H}]$), and α metallicity. The α metallicity is used to specify the abundance of the elements O, Ne, Mg, Si, S, Ar, Ca, and Ti relative to Fe. However, we found no data on α metallicity for the target stars, so we treated each as having the solar value. For T_{eff} , $\log_{10} g$, and $[\text{Fe}/\text{H}]$, we found literature values for all stars, as listed in Table 2.

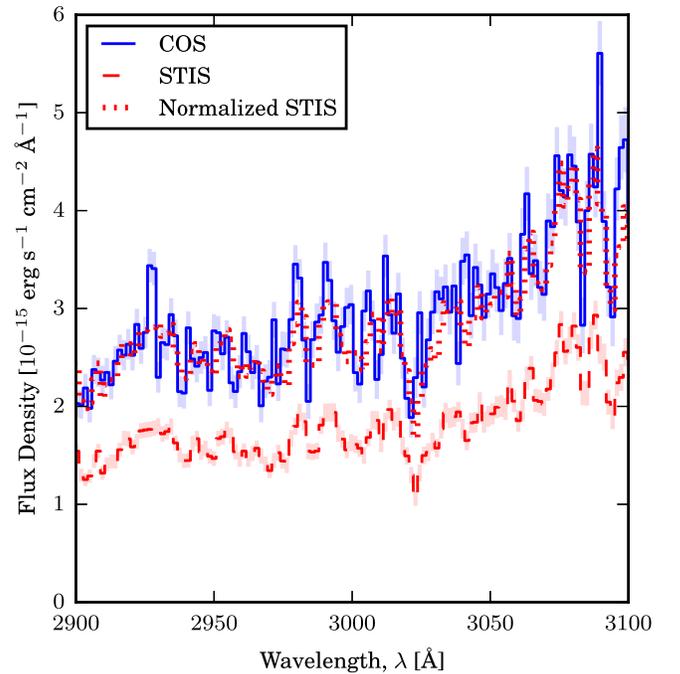


Figure 3. Example of normalizing STIS to COS data. The figure shows a portion of the NUV data for GJ 436 from the STIS G230L and COS G230L observations, identically binned for comparison. The COS flux exceeds the STIS flux by a wide margin. Normalizing the STIS data in lieu of possible slit losses produces good agreement.

The PHOENIX model output provided by Husser et al. (2013) has an arbitrary scale. Thus, the output must be normalized to match the absolute flux level of the star. To constrain this normalization using data on the star, we collected all external photometry for the star returned by the VizieR Photometry Viewer¹⁴ within a $10''$ search radius. An exception was GJ 667C, where companion stars A and B prevented a

¹⁴ <http://vizier.u-strasbg.fr/vizier/sed/>

Table 2
Selected Properties of the Stars in The Sample

Star	Type	d (pc)	V	Ref	MUSCLES T_{eff} (K)	Literature T_{eff} (K)	Ref	$\log g$ (cm s^{-2})	Ref	[Fe/H]	Ref
GJ 1214	M4.5	14.6	14.68 ± 0.02	(1)	2935 ± 100	2817 ± 110	(2)	5.06 ± 0.52	(3)	0.05 ± 0.09	(2)
GJ 876	M5	4.7	10.192 ± 0.002	(4)	3062^{+120}_{-130}	3129 ± 19	(5)	4.93 ± 0.22	(3)	0.14 ± 0.09	(2)
GJ 436	M3.5	10.1	10.59 ± 0.08	(6)	3281 ± 110	3416^{+54}_{-61}	(7)	4.84 ± 0.16	(3)	-0.03 ± 0.09	(2)
GJ 581	M5	6.2	10.61 ± 0.08	(6)	3295 ± 140	3442 ± 54	(8)	4.96 ± 0.25	(3)	-0.20 ± 0.09	(2)
GJ 667C	M1.5	6.8	10.2	(9)	3327 ± 120	3445 ± 110	(2)	4.96 ± 0.25	(3)	-0.50 ± 0.09	(2)
GJ 176	M2.5	9.3	10.0	(10)	3416 ± 100	3679 ± 77	(5)	4.79 ± 0.13	(3)	-0.01 ± 0.09	(2)
GJ 832	M1.5	5.0	8.7	(10)	3816 ± 250	3416 ± 50	(11)	4.83 ± 0.15	(3)	-0.17 ± 0.09	(2)
HD 85512	K6	11.2	7.7	(10)	4305^{+120}_{-110}	4400 ± 45	(12)	4.4 ± 0.1	(12)	-0.26 ± 0.14	(12)
HD 40307	K2.5	13.0	7.1	(10)	4783 ± 110	4783 ± 77	(12)	4.42 ± 0.16	(12)	-0.36 ± 0.02	(12)
HD 97658	K1	21.1	7.7	(10)	5156 ± 100	5170 ± 50	(13)	4.65 ± 0.06	(14)	-0.26 ± 0.03	(14)
ϵ Eri	K2.0	3.2	3.7	(15)	5162 ± 100	5049 ± 48	(12)	4.45 ± 0.09	(12)	-0.15 ± 0.03	(12)

Note. We located many of these parameters through the PASTEL (Soubiran et al. 2010) catalog, but have provided primary references in this table.

References. (1) Weis (1996), (2) Neves et al. (2014), (3) Santos et al. (2013), (4) Landolt (2009), (5) von Braun et al. (2014), (6) Høg et al. (2000), (7) von Braun et al. (2012), (8) Boyajian et al. (2012), (9) Mermilliod (1986), (10) Koen et al. (2010), (11) Houdebine (2010), (12) Tsantaki et al. (2013), (13) Van Grootel et al. (2014), (14) Valenti & Fischer (2005), (15) Ducati (2002).

position-based search. Instead, we collected photometry specifically associated with the object from the Denis catalog, the UCAC4 catalog, and the HARPS survey. For all other stars, we verified that no other sources fell within the search radius in 2MASS imagery. References for the photometry we collected are given in Table 3.

For several stars, simply retrieving a PHOENIX spectrum from the Husser et al. (2013) grid using stellar parameters found in the literature produced a spectrum with an overall shape that did not match the collected photometry. The overall spectral shape is primarily driven by T_{eff} , so, to correct this mismatch, we wrote an algorithm to search the Husser et al. (2013) grid of PHOENIX spectra for the best-fit T_{eff} . Our fitting algorithm operated by taking a supplied T_{eff} and the literature values for $\log_{10} g$ and [Fe/H] and tri-linearly interpolating a spectrum from the Husser et al. (2013) grid. It then computed the best-fit normalization factor that matched the spectrum to stellar photometry. The normalization factor was computed analytically via a min- χ^2 fit to the photometry under the assumption of identical signal-to-noise ratio (S/N) for each point. We estimated the S/N as the rms of the normalized residuals, that is

$$\frac{F_{o,i}}{\sigma_i} = \left[\frac{1}{N} \sum_i \frac{(F_{o,i} - F_{c,i})^2}{F_{o,i}^2} \right]^{-1/2}, \quad (1)$$

where σ_i is the uncertainty on the observed flux $F_{o,i}$, $F_{c,i}$ is the synthetic flux computed by applying the transmission curve of the filter used to measure $F_{o,i}$ to the PHOENIX model, and i indexes the N available flux measurements. Estimating measurement uncertainties permits the inclusion of data lacking quoted uncertainties and mitigates possible underestimation of uncertainties in those data for which they are given. Uncertainties on the best-fit normalization factors span 0.007% (ϵ Eri) to 6% (GJ 1214) with a median uncertainty of 0.4%, and the uncertainty in normalization correlates well with the target V magnitude.

After normalization, the algorithm checked for outlying photometry by computing the deviation for which the false alarm probability of a point occurring beyond the deviation was <10%. Points beyond that deviation were culled and the fit

Table 3
References for Stellar Photometric Measurements

Star	Photometry References
GJ 1214	(1), (2), (3), (4), (5), (6), (7), (8), (9)
GJ 876	(10), (11), (12), (7), (13), (6), (14), (15), (16), (3), (17), (18), (19), (20), (9)
GJ 436	(21), (22), (12), (2), (7), (13), (14), (3), (17), (6), (19), (23), (9)
GJ 581	(10), (21), (1), (24), (7), (13), (17), (12), (14), (19), (23), (9)
GJ 667C	(25), (26), (27)
GJ 176	(11), (22), (7), (13), (6), (14), (15), (16), (12), (18), (19), (9)
GJ 832	(28), (10), (29), (30), (7), (13), (14), (16), (31), (32), (23), (19), (33), (8), (9)
HD 85512	(28), (34), (35), (36), (37), (30), (23), (13), (14), (31), (18), (32), (33), (8), (9)
HD 40307	(28), (34), (35), (11), (37), (30), (23), (13), (9), (3), (18), (32), (33), (8), (38)
HD 97658	(11), (37), (30), (3), (13), (39), (31), (18), (32), (33), (8), (9)
ϵ Eri	(35), (37), (30), (40), (41), (13), (14), (42), (9), (18), (32), (33), (8), (43)

References. (1) Zacharias et al. (2004a), (2) Triaud et al. (2014), (3) Santos et al. (2013), (4) Wright et al. (2011), (5) Cutri et al. (2014), (6) Rojas-Ayala et al. (2012), (7) Lépine & Gaidos (2011), (8) Cutri et al. (2012), (9) Cutri et al. (2003), (10) Winters et al. (2015), (11) Ammons et al. (2006), (12) de Bruijne & Eilers (2012), (13) Röser et al. (2008), (14) Salim & Gould (2003), (15) Zacharias et al. (2004b), (16) Bonfils et al. (2013), (17) Jenkins et al. (2009), (18) Pickles & Depagne (2010), (19) Gaidos et al. (2014), (20) Finch et al. (2014), (21) Roeser et al. (2010), (22) Reid et al. (2004), (23) Abrahamyan et al. (2015), (24) Bryden et al. (2009), (25) The DENIS Consortium (2005), (26) Delfosse et al. (2013), (27) Zacharias et al. (2012), (28) Girard et al. (2011), (29) Neves et al. (2013), (30) van Leeuwen (2007), (31) Myers et al. (2015), (32) Anderson & Francis (2012), (33) Kharchenko (2001), (34) Lawler et al. (2009), (35) Holmberg et al. (2009), (36) Anglada-Escudé & Butler (2012), (37) Soubiran (2010), (38) Eiroa et al. (2013), (39) Bailer-Jones (2011), (40) Fabricius et al. (2002), (41) Haakonsen & Rutledge (2009), (42) Aparicio Villegas et al. (2010), (43) Gould (1879).

recomputed to convergence. Once converged, the algorithm computed the likelihood of the model given the data. This then permitted a numerical search for the maximum-likelihood T_{eff} .

Once the best-fit T_{eff} was found, the algorithm sampled the likelihood function to find the 68.3% confidence interval on

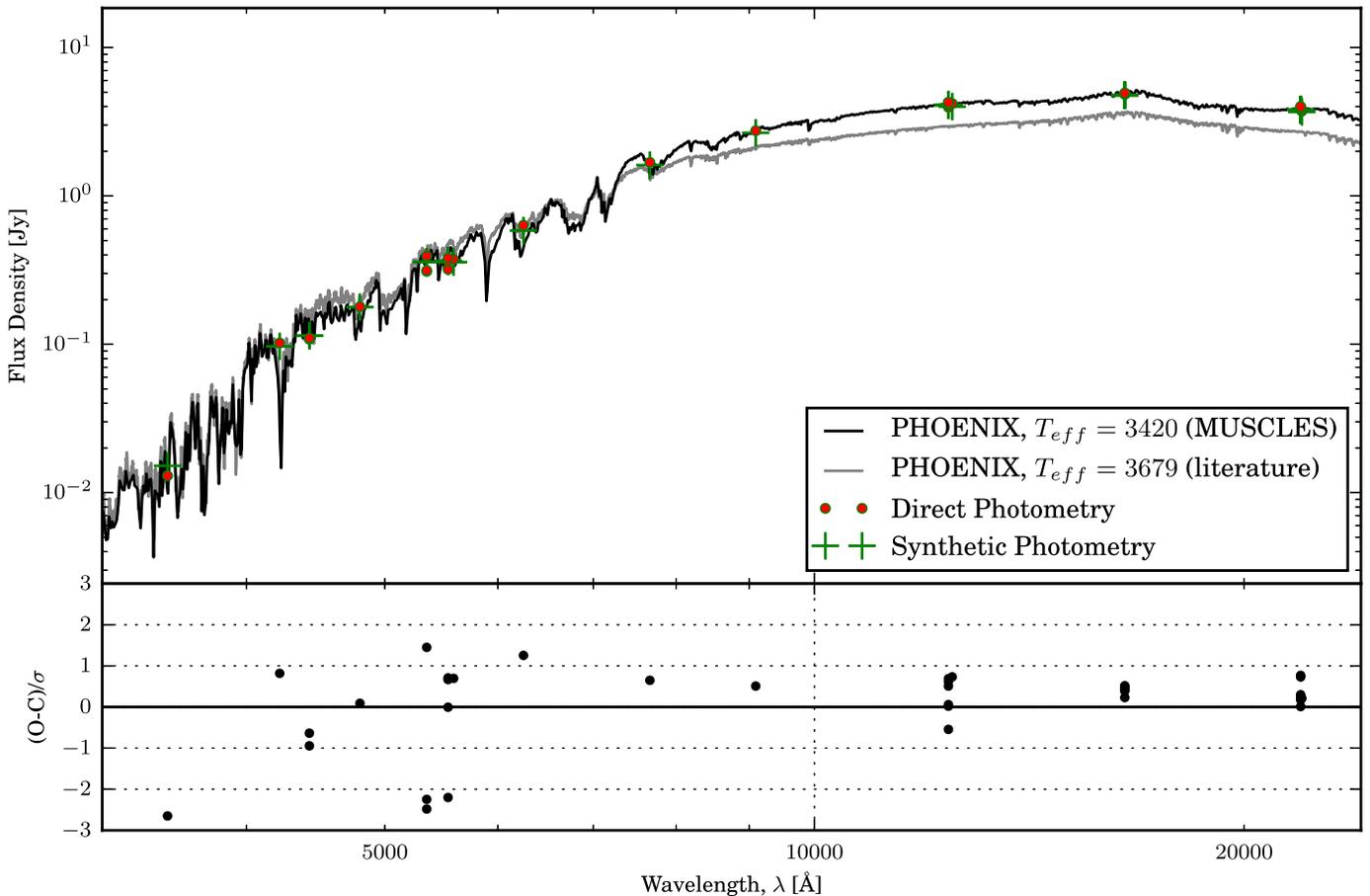


Figure 4. Normalization and T_{eff} fit of PHOENIX model output to non-MUSCLES photometry for GJ 176. Synthetic photometry (green crosses) computed from the PHOENIX spectrum with best-fit T_{eff} (black line) agrees well with all photometric data (red dots). In contrast, the PHOENIX spectrum where T_{eff} is fixed to the literature value (gray line) produces a noticeably poor fit to the photometry. Photometric data are plotted at the mean filter wavelength.

T_{eff} . For this search, the photometry was fixed to the outlier-culled list and associated S/N estimate from the best fit. This confidence interval provides a statistical uncertainty estimate, but the assumptions made in the fitting process (PHOENIX model, constant S/N estimated from residuals) introduces a further, systematic uncertainty. We estimated this systematic uncertainty to be 100 K and added it in quadrature to the statistical uncertainty to produce a final uncertainty estimate for the T_{eff} values we computed for each star.

Our best-fit T_{eff} values and corresponding literature values are listed in Table 2. Figure 4 illustrates the discrepancy between the literature T_{eff} values and the photometry for the worst case. The figure includes one fit computed with T_{eff} as a free parameter and another with T_{eff} fixed to the literature value. The shape of the PHOENIX spectrum interpolated at the literature value does not conform to the photometry.

After finding the best-fit PHOENIX spectrum, we used it to normalize the STIS G430L data according to the ratio of integrated fluxes in the overlap redward of 3500 Å.

2.5. Creating Composite Spectra

We spliced together the spectra from each source to create the final spectrum. When splicing, observations were always (with a few exceptions mentioned in the next section) given preference over the APEC or PHOENIX models. When two observations were spliced, we chose the splice wavelength to

minimize the error on the integrated flux in the overlapping region. This causes the splice locations to vary somewhat from star to star (see Figure 2). Once all of the spectra were joined, we filled the remaining small gaps that resulted from removing telluric lines (along with a slight separation between the E230M and E230H spectra in ϵ Eri) with a quadratic fit to the continuum in a region 20 times the width of the gap. Throughout the process, we propagated exposure times, start date of first exposure, end date of last exposure, normalization factors, data sources, and data quality flags on a pixel-by-pixel basis into the final data product.

2.6. Special Cases

The diversity of targets and data sources included in this project necessitated individual attention in the data reduction process. In all cases, spectra were separately examined and suspect data, particularly data near the detector edges, were culled. In a variety of cases, detailed below, we tweaked the data reduction process.

The *HST* STIS G430L data for all stars show large scatter with many negative flux bins at the short-wavelength end of the spectrum. For most stars, this region is small enough to be significantly overlapped by COS G230L data that we have used in the composite SED. However, for GJ 667C and GJ 1214 this region is much larger than the COS G230L overlap. In these cases, we culled the STIS G430L data from the long-

wavelength edge of the COS G230L spectra to $\sim 3850 \text{ \AA}$, roughly where the negative-flux pixels of the STIS G430L data cease, and filled the gap with the PHOENIX model. While the PHOENIX models used here omit emission from the stellar upper atmosphere that dominates short-wavelength flux (see Section 3.4), this omission does not begin to have a substantial effect until shortward of where the COS G230L coverage begins in other targets.

GJ 667C showed fluxes for all STIS data several times below that of the COS data. Inspection of the acquisition images revealed the spectrograph slit was very poorly centered on the source. We did not alter the reduction process for this star, but we note that the STIS data were normalized upward by large factors in this case. Specifically, we computed normalization factors of 3.5 for the STIS G140M data, 5.1 for the STIS G230L data, and 4.2 for the STIS G430L data.

The STIS G140M spectra of GJ 1214, GJ 832, GJ 581, and GJ 436 and the STIS G230L spectrum of GJ 1214 were improperly extracted by the CalSTIS pipeline because the pipeline algorithm could not locate the spectral trace on the detector. However, the spectral data were present in the fluxed two-dimensional (2D) image of the spectrum (the .x2d files). We extracted spectra for these targets by summing along the spatial axis in a region of pixels centered on the signal in the spectral images with a proper background subtraction and correction for excluded portions of the PSF. We checked our method using stars where CalSTIS succeeded in locating and extracting the spectrum and found good agreement.

The *Chandra* data for GJ 1214 provided only an upper limit on the X-ray flux. As a result, we decided to use a previous model fit to *XMM-Newton* data (Lalitha et al. 2014) to fill the X-ray portion of GJ 1214’s spectrum.

We did not acquire X-ray data for HD 97658. To fill this region of the spectrum, we use data from HD 85512 scaled by the ratio of the bolometric fluxes of the two stars. These stars have nearly identical Fe XII $\lambda 1242$ emission relative to their bolometric flux. The Fe XII ion has a peak formation temperature above 10^6 K (Dere et al. 2009), associating it with coronal emission. Thus, the similarity in Fe XII emission between HD 97658 and HD 85512 suggests similar levels of coronal activity. This conclusion is further supported by the similar ages estimated by Bonfanti et al. (2016) of $9.70 \pm 2.8 \text{ Gyr}$ for HD 97658 and $8.2 \pm 3.0 \text{ Gyr}$ for HD 85512.

Both GJ 581 and GJ 876 were observed by *Chandra* at two markedly different levels of X-ray activity. We include in the panchromatic SEDs the observations taken when the stars were less active.

Finally, the spectral image of the STIS G230L exposure of GJ 436 shows a faint secondary spectrum separated by about 80 mas from the primary. It is very similar in spectral character, if not identical, to the GJ 436 spectrum. An identical exposure from 2012 does not show the same feature, but this is unsurprising given the high proper motion of GJ 436 (van Leeuwen 2007). If a second source is a significant contributor, this could impart an upward bias to the GJ 436 flux.

2.7. Notes on Data Quality

Users should keep several important characteristics of the panchromatic SEDs in mind when using them.

2.7.1. Flares

These stars exhibited flares in the UV, some very large (see Paper I), during the observations. The rates of such flares on these stars is mostly unconstrained: there is little or no previous data that could provide a good estimate of whether the observed flares were typical or atypical for the target. Thus, the safest assumption is to treat the MUSCLES observations as typical. As such, we did not attempt to separate the data into times of flare and quiescence. These observations should thus be treated as roughly typical of any average of one to a few hours of UV data for a given star. These flares will be the subject of a future publication (R. O. P. Loyd et al. 2016b, in preparation).

2.7.2. Wavelength Calibration

We observed some mismatches in the wavelengths of spectral features in the NUV data from COS and STIS. This mismatch is most pronounced at the Si II $\lambda\lambda 1808, 1816$ lines, where the line centers in the STIS E230M data for ϵ Eri (the only STIS observation that resolves the lines) match the true line wavelengths, but the COS G230L data of all other targets show the lines shifted $\sim 4 \text{ \AA}$ ($\sim 660 \text{ km s}^{-1}$) blueward. This shift in the COS G230L wavelength solution is not present at the next recognizable spectral feature redward of Si II, Mg II. Flux from Si II and Mg II is captured on different spectral “stripes” on the COS G230L detector, suggesting that the entire stripe capturing Si II flux might be poorly calibrated in comparison to the stripe capturing Mg II flux. The Mg II $\lambda\lambda 2796, 2803$ lines are in an area of overlap between the COS G230L data and STIS E230M (K stars) and G230L (M stars) data, enabling a direct comparison of the wavelength solutions for these spectra at Mg II. Making this comparison reveals that STIS E230H and COS G230L data agree to within 1 \AA , while STIS G230L and COS G230L data can disagree by up to 2 \AA ($\sim 330 \text{ km s}^{-1}$). In the latter cases, the COS data more closely match the true wavelength of the Mg II lines.

The NUV discrepancy for the narrow-slit STIS G230L observations at Mg II could result from the same target centering issues that may cause the flux discrepancies discussed in Section 2.3.1. We do not understand the source of the erroneous COS wavelength solution at the Si II lines.

In the FUV, wavelength calibrations between COS G130M and STIS G140M data at the N V $\lambda\lambda 1238, 1242$ lines typically agree to within a single pixel width of the STIS G140M data ($\sim 0.05 \text{ \AA}$, 10 km s^{-1}).

We did not alter the wavelength calibration of any spectra. Neither did we attempt to shift the spectra to the rest frame of the star. Such a shift would be $< 0.05\%$ for any target ($< 0.5 \text{ \AA}$ at 1000 \AA). To place the wavelength miscalibrations and target redshifts in context, we note that the molecular cross section spectra used in Section 3.3 have a resolution of 1 \AA .

2.7.3. Negative Flux

Users will notice bins with negative flux density in many low-flux regions in the UV. This is a result of subtracting a smoothed background count rate from regions where noise dominates the signal (Section 3.4¹⁵ of the STIS Data Handbook, Bostroem & Proffitt 2011 and Section 3.4¹⁶ of the COS

¹⁵ http://www.stsci.edu/hst/stis/documents/handbooks/currentDHB/ch3_stis_calib5.html

¹⁶ http://www.stsci.edu/hst/cos/documents/handbooks/datahandbook/ch3_cos_calib5.html

Data Handbook, Fox et al. 2015). While a negative flux is unphysical, the background subtraction serves to produce an unbiased estimate of the flux when low-flux regions are integrated. Thus, we left the negative bins in the low-flux regions unaltered.

2.7.4. The Rayleigh–Jeans Tail

The truncation of the PHOENIX spectra at $5.5 \mu\text{m}$ results in the omission of flux that contributes a few percent to the bolometric flux of the stars. The bolometric flux estimate determines the relative level at which each wavelength regime contributes to the stellar SED, so accuracy is important. We therefore compute and include in the data product a bolometric flux value for each target that incorporates the integral of a blackbody fit to the PHOENIX spectrum from the red end of the PHOENIX range at $5.5 \mu\text{m}$ to ∞ . Whenever we present bolometrically normalized fluxes, we use this more accurate value as opposed to the integral of the MUSCLES SED alone.

2.7.5. ϵ Eri STIS E230M/E230H Data Compared with PHOENIX Output

Unlike the other targets, ϵ Eri was too bright to permit the collection of *HST* COS observations in the NUV, so only *HST* STIS data, specifically using the E230M and E230H gratings, were collected. Because of the lack of COS data that was used to correct systematically lower flux levels in the STIS NUV data of other stars (Section 2.3.1), we examined the ϵ Eri STIS NUV data closely. This star has a high enough T_{eff} for photospheric flux to contribute significantly in the NUV. Thus, the PHOENIX spectrum for ϵ Eri can be meaningfully compared to the observation data in this regime.

The lower envelope of the PHOENIX spectrum matches very well with that of the E230M and E230H data. However, some portions of the PHOENIX spectrum show emission features well above that seen in the E230M and E230H spectra. This results in the integrated flux of the PHOENIX spectrum exceeding that of the observations by 28% for E230M and 43% for E230H. Because of the agreement in the lower envelopes of the spectra, we conclude that this mismatch is likely caused by inaccuracy in the PHOENIX spectrum rather than in the observations. At longer wavelengths, the *HST* STIS G430M observation of ϵ Eri covering $\sim 3800\text{--}4100 \text{ \AA}$ agrees with the PHOENIX spectrum to 0.2%, suggesting the accuracy of both the observations and PHOENIX output is good in that regime.

2.7.6. Transits

We did not attempt to avoid planet transits during observations of the MUSCLES targets to facilitate the observatory scheduling process. As a result, some observations overlap with planet transits. We checked for overlap by acquiring transit ephemerides for all hosts where these ephemerides were well established from the NASA Transit and Ephemeris Service¹⁷ and comparing these in-transit time ranges to the time ranges of *HST* observations.

GJ 1214b transited during one of the *HST* observations: one of the three COS G160M exposures is almost fully within GJ 1214b’s geometric transit. However, the $\sim 1\%$ transit depth of

GJ 1214b (Carter et al. 2011), is insignificant in comparison to the 34% uncertainty on the integrated G160M flux.

GJ 436b was undergoing geometric transit ingress at the end of the third COS G130M exposure and transit egress at the start of the fourth, according to the Knutson et al. (2011) ephemeris. The G130M exposures were sequential, broken up only by Earth occultations. Consequently, the last four of the five G130M exposures fall within the range of the extended Ly α transit that begins two hours prior and lasts at least three hours following the geometric transit and absorbs 56% of the stellar Ly α emission (Kulow et al. 2014; Ehrenreich et al. 2015). Thus, these observations should be treated as lower limits to the out-of-transit emission of GJ 436 from ions that might be present in the planet’s extended escaping cloud. The geometric transit depth of GJ 436 is under 1% (Torres et al. 2008), so only an extended cloud of ions could effect the G130M spectra by an amount that is significant relative to uncertainties. We inspected the G130M data for evidence of transit absorption by measuring the flux of the strongest emission lines as a function of time and found no such evidence. However, because all exposures may be affected by an extended cloud, the lack of a clear transit dip might not be conclusive. We intend to explore the COS G130M transit data in greater depth in a future work (R. O. P. Loyd et al. 2016a, in preparation). The reconstructed Ly α flux stitched into the panchromatic SEDs is not affected. It was created from separate STIS observations that occurred outside of transit.

3. DISCUSSION

We display the primary data product of the MUSCLES survey, the panchromatic SEDs for each target, in Figures 5 and 6. We also include a spectrum of the Sun for comparison to the MUSCLES stars. The solar spectrum we present in these figures and in the discussion following is that assembled by Woods et al. (2009) from several contemporaneous datasets covering a few days of “moderately active” solar emission. We adopt a value of $1.361 \times 10^6 \text{ erg s}^{-1} \text{ cm}^{-2}$ for the bolometric solar flux at Earth, i.e., Earth’s insolation, per resolution B3 of IAU General Assembly XXIX (2015).

For a detailed discussion of emission line strengths and the correlation of fluxes from lines with different formation temperatures, see Paper I. Here we will examine the UV continuum, the degree of deviation from purely photospheric models, and photodissociation rates on orbiting planets.

3.1. The FUV Continuum

Hitherto, the sensitivities of FUV instruments, such as the *International Ultraviolet Explorer* and *HST* STIS, have rendered observations of anything but isolated emission lines in spectra from low-mass stars challenging. The greater effective area and lower background rate of *HST* COS facilitates observations of less prominent spectral features. Linsky et al. (2012) used this advantage to study the FUV continua of solar-mass stars, finding that more rapidly rotating stars showed higher levels of FUV surface flux and corresponding brightness temperature.

In a similar vein, we searched for evidence of an FUV continuum in the spectra of the MUSCLES stars. Continuum regions were identified by eye through careful examination of the spectra of all 11 targets, similar to the methodology of France et al. (2014), and typically range from 0.7 to 1 \AA in width. The

¹⁷ <http://exoplanetarchive.ipac.caltech.edu/cgi-bin/TransitView/nph-visibletbls?dataset=transits>

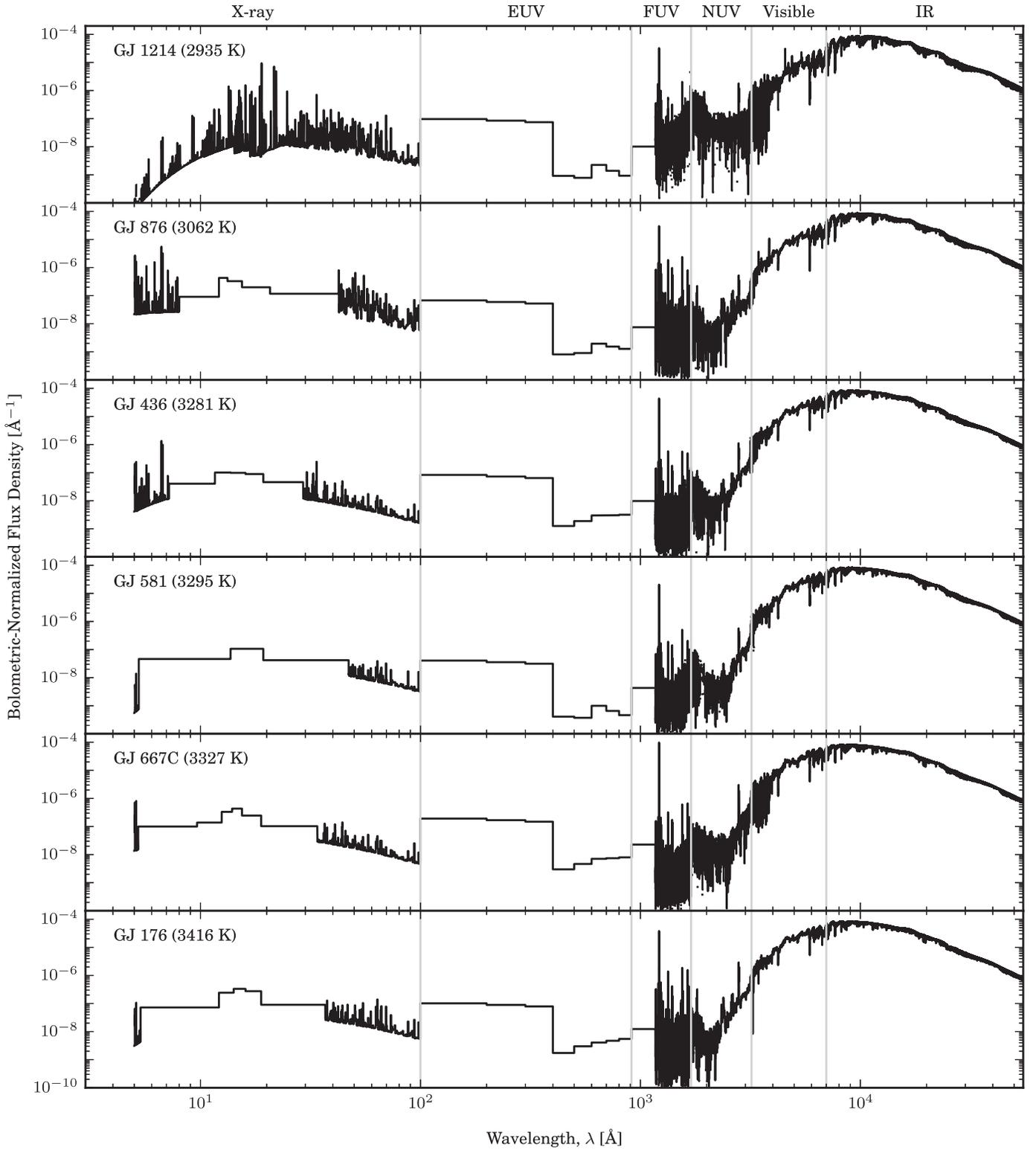


Figure 5. Spectral energy distributions of the MUSCLES stars, normalized by their bolometric flux such that they integrate to unity. Axes have identical ranges to facilitate comparisons, each spanning 3–55000 Å in wavelength and 10^{-10} – 2×10^{-4} Å $^{-1}$ in flux density. Light gray vertical lines show the division between the spectral regions labeled at the top of the plot. The best-fit effective temperature of each star computed via our PHOENIX model grid search (Section 2.4) is listed in parenthesis next to its name. The normalization enables easy scaling to any star–planet distance; e.g., for an Earth-like planet, multiply by Earth’s insolation, 1361 W m^{-2} .

selected regions constitute 157.9 \AA of continuum spanning 1307–1700 Å. These regions have a bandwidth-weighted mean wavelength of 1474.3 Å. We integrated the flux in each of these bands to create continuum spectra for the targets.

The continuum spectrum of the brightest target, ϵ Eri, is displayed in Figure 7. This continuum shows a shape that strongly suggests a recombination edge occurring between 1500 and 1550 Å. This is consistent with the $\sim 1521 \text{ \AA}$ limit of

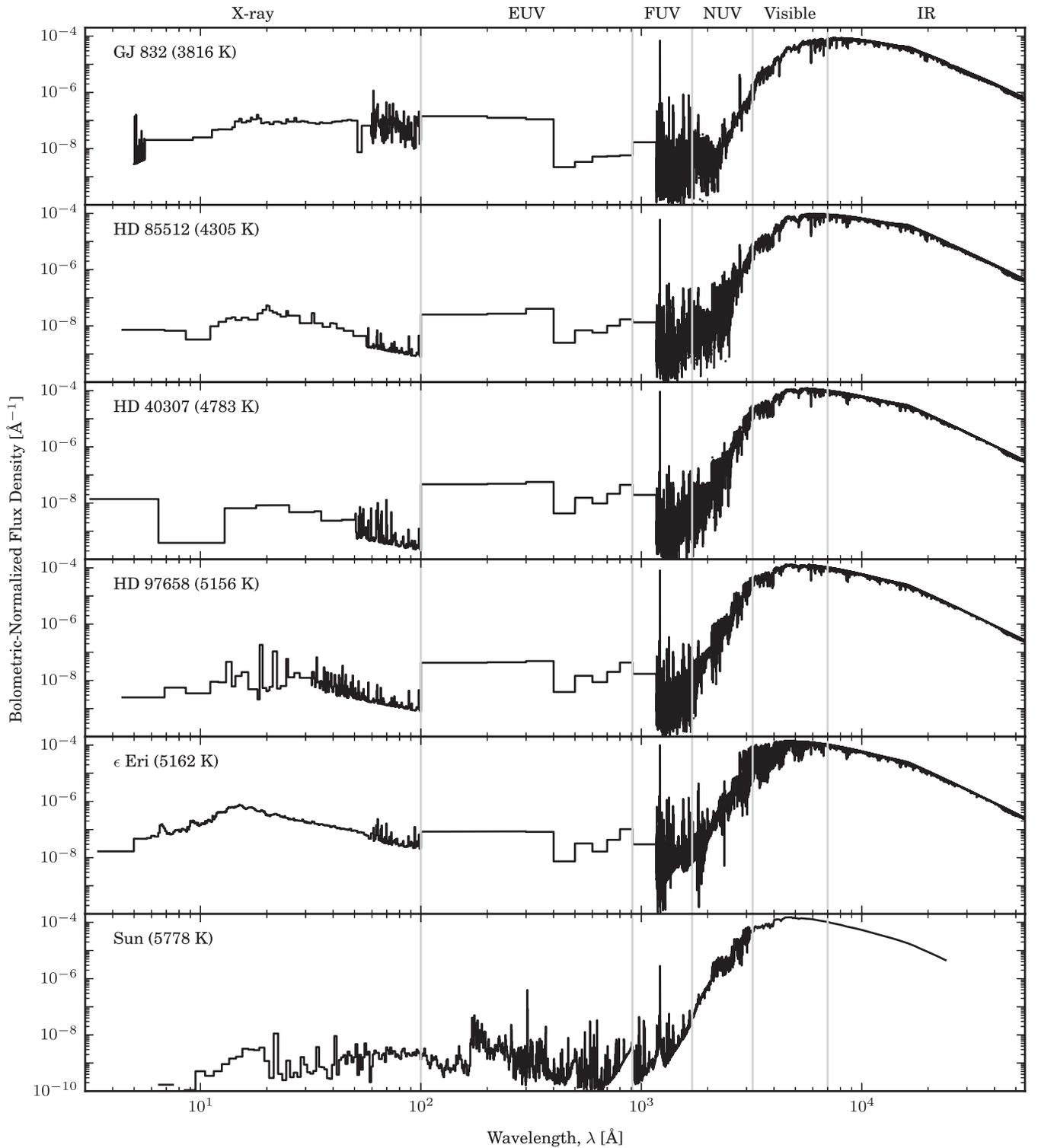


Figure 6. Spectral energy distributions of the MUSCLES stars, continued from Figure 5.

Si II recombination to Si I. Models of the solar chromosphere have predicted this edge, but it is not observed in the solar FUV spectrum (Fontenla et al. 2009), nor was it observed in the continua of solar-mass stars studied by Linsky et al. (2012).

Data from other targets did not have sufficient S/N to show a clear continuum shape. However, by integrating the flux in all 157.9 Å of continuum bands, some level of continuum was detected at $>3\sigma$ significance for 5/11 targets. Note that the

negative values in Table 4 are not concerning given their large error bars. A simple detection of continuum emission does not indicate whether this emission is significant relative to line emission in the FUV. To quantify this significance, we compute the fractional contribution of the flux in the continuum bands to the total FUV flux over the full range containing the continuum bands (1307–1700 Å) and present the results in the last column of Table 4. This fraction does not account for the

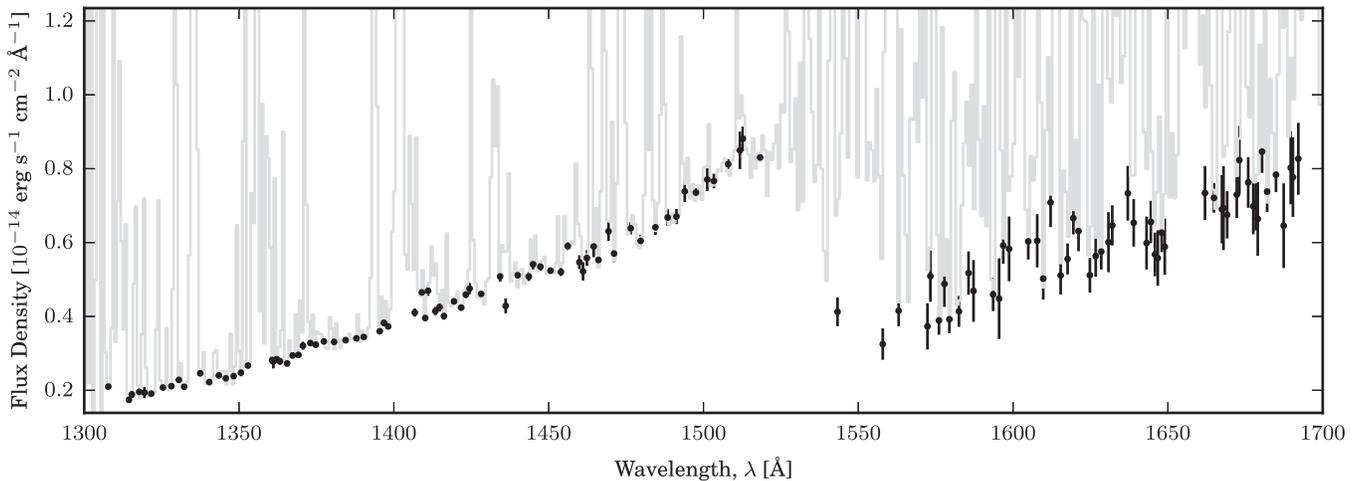


Figure 7. Continuum of ϵ Eri, the brightest source with the clearest continuum. Each black point is the average flux density in a continuum band. The width of each band is roughly the size of the point or smaller. The full spectrum (i.e., non-continuum) is plotted in gray in the background, rebinned to $R = 2000$ for display. The edge occurring between 1500 and 1550 Å is consistent with recombination of Si II to Si I at ~ 1521 Å.

continuum flux present within emission line regions. Therefore, it serves as a lower limit on the total contribution of continuum within the 1307–1700 Å range. This contribution is at least on the order of 10% for the stars where it is detected, and where it is not detected the uncertainties do not rule out similar levels.

3.2. Estimating SEDs for Stars Not in the MUSCLES Catalog

A means of estimating the UV SED for any star without *HST* observations would be very useful. We investigated the potential for using photometry in the *GALEX* FUV and NUV bands to empirically predict fluxes in other UV bands. However, the *GALEX* All-sky Survey (AIS; Bianchi et al. 2011) contains magnitudes for only half of the MUSCLES sample. The MUSCLES stars represent the brightest known M and K dwarf hosts. Therefore, they will be better represented than a volume-limited sample of such stars in the *GALEX* AIS catalog. That half do not appear in the catalog suggests the *GALEX* AIS catalog, while an excellent tool for population studies with few restrictions on sample selection, is not a generally useful tool for studies involving one or a few preselected targets. Nonetheless, we checked for a correlation of *GALEX* FUV magnitudes for the six targets in the AIS catalog with their integrated FUV flux from *HST* data and found no correlation. It bears mentioning that while most M and K dwarf planet hosts may not appear in *GALEX* catalogs, the *GALEX* survey data have proved useful both for prospecting for low-mass stars (Shkolnik et al. 2011) and population studies of these stars (Shkolnik & Barman 2014; Jones & West 2016).

There are other means, besides *GALEX* data, of estimating broadband and line fluxes for a low-mass star without complete knowledge of its SED. To this end, Paper I provides useful relations between broadband fluxes and emission line fluxes. Paper I also quantifies levels of FUV activity, taken as the ratio of the integrated FUV flux to the bolometric flux, and shows that they are roughly constant. This consistency suggests that the median ratio of FUV to bolometric flux of the MUSCLES sample would be appropriate as a first-order estimate of that same ratio for any similar low-mass star. Further, the MUSCLES spectra, normalized by their bolometric flux, can serve as rough template spectra for a low-mass star with no

Table 4
FUV Continuum Measurements

Star	FUV Continuum Flux ($\text{erg s}^{-1} \text{cm}^{-2} \text{Å}^{-1}$)	Detection Significance σ	Fraction of Flux ^a
GJ 1214	$-1.2 \pm 1.1 \times 10^{-15}$...	-0.753
GJ 876	$6.7 \pm 1.2 \times 10^{-15}$	5.7	0.080
GJ 436	$1.1 \pm 1.4 \times 10^{-15}$...	0.082
GJ 581	$-8 \pm 13 \times 10^{-16}$...	-0.201
GJ 667C	$1.1 \pm 1.1 \times 10^{-15}$...	0.088
GJ 176	$4.1 \pm 1.2 \times 10^{-15}$	3.5	0.070
GJ 832	$6.4 \pm 1.1 \times 10^{-15}$	6.1	0.122
HD 85512	$6.4 \pm 5.4 \times 10^{-15}$...	0.114
HD 40307	$8.2 \pm 4.8 \times 10^{-15}$...	0.147
HD 97658	$1.13 \pm 0.18 \times 10^{-14}$	6.3	0.201
ϵ Eri	$8.128 \pm 0.038 \times 10^{-13}$	215.9	0.128

Note.

^a Integral of flux within continuum bands divided by the integral of all flux over the range containing those bands.

available spectral observations. If the star’s bolometric flux is known, this can be multiplied with the template normalized spectrum to provide absolute flux estimates.

However, age is likely a factor in the consistency in FUV activity among MUSCLES targets. Age estimates are not presently available for the full MUSCLES sample; however, ages of the MUSCLES stars are likely to be consistent with other nearby field stars of similarly weak Ca II H and K flux (see Paper I). These have ages of several Gyr (Mamajek & Hillenbrand 2008). Stars with ages under a Gyr are likely to have higher levels of FUV and NUV activity compared to the MUSCLES sample. This conclusion follows from the recent work of Shkolnik & Barman (2014) that demonstrated a dependence of excess (i.e., non-photospheric) FUV and NUV flux on stellar age for early M dwarfs. Shkolnik & Barman (2014) found that the excess FUV and NUV flux of their stellar samples remained roughly constant at a saturated level until an age of a few hundred Myr. After this age, the excess flux declines by over an order of magnitude as the stars reach ages of several Gyr. Because of this age dependence in FUV flux,

the consistency in FUV activity of the MUSCLES stars should not be taken to be representative of stars with ages under a Gyr.

3.3. Photodissociation

The composition of a planetary atmosphere depends on a complex array of factors including mass transport, geologic sources and sinks, biological activity, impacts, aqueous chemistry, stellar wind, and incident radiation (e.g., Matsui & Abe 1986; Lammer et al. 2007; Seager & Deming 2010; Hu et al. 2012; Kaltenecker et al. 2013). Many of these factors require assumptions weakly constrained by what is known of solar system planets. However, the MUSCLES data directly characterize the radiation field incident upon planets around the 11 host stars. This allows the use of MUSCLES data to address two top-level questions regarding atmospheric photochemistry without detailed modeling of specific atmospheres: (1) What is the relative importance of various spectral features to the photodissociation of important molecules, and (2) how does this differ between stars?

To quantitatively answer these questions, we examined the photodissociation rates of various molecules as a function of wavelength resulting from the radiative input of differing stars. Because we leave atmospheric modeling to future work, we did not introduce any attenuation of the stellar SED by intervening material. In other words, we assumed direct exposure of the molecules to the stellar flux. However, for wavelengths shortward of the ionization threshold of the molecule, we assumed that all photon absorptions result in ionization rather than dissociation. Consequently, for each molecule we ignore stellar flux shortward of that molecule’s ionization threshold. Because we assumed direct exposure to the stellar flux, we refer to the photodissociation rates we have computed as “unshielded.”

3.3.1. Method of Computing “Spectral” Photodissociation Rates

To determine “spectral” photodissociation rates, $j(\lambda)$, we multiplied the wavelength-dependent cross sections, $\sigma(\lambda)$, by the photon flux density, $F_p(\lambda)$, and the sum of the quantum yields for all dissociation pathways, $q(\lambda)$. The quantum yield of a pathway expresses the probability that the molecule will dissociate through that pathway if it absorbs a photon with wavelength λ . To wit,

$$j(\lambda) = F_p(\lambda)\sigma(\lambda)\sum_i q_i(\lambda). \quad (2)$$

Note that the stellar spectrum must be specified in *photon* flux density ($\text{s}^{-1}\text{cm}^{-2}\text{\AA}^{-1}$) rather than energy flux density ($\text{erg s}^{-1}\text{cm}^{-2}\text{\AA}^{-1}$) ($F_p = F/(hc/\lambda)$).

The photodissociation cross sections and quantum yields we used come from the data gathered by Hu et al. (2012) and presented in their Table 2, updated to include the high-resolution cross section measurements of O_2 in the Schumann–Runge bands from Yoshino et al. (1992). We extrapolated the H_2O cross section data, as is conventional, from $\sim 2000\text{\AA}$ to the dissociation limit at 2400\AA using a power-law fit to the final two available data points. The cross sections only apply to photodissociation from the ground state. We also do not include H_2 dissociation by excitation of the Lyman and Werner bands. Physically, j in Equation (2) represents the rate at which unshielded molecules are photodissociated by stellar photons per spectral element at λ . Thus the units of j are $\text{s}^{-1}\text{\AA}^{-1}$.

The absolute level of j is tied to a specific distance from the star since F_p drops with the inverse square of distance. For the values we present in the remainder of Section 3.3, we set the F_p level of each star’s spectrum such that the bolometric energy flux (“instellation”), $I = \int_0^\infty F(\lambda)d\lambda$, was equivalent to Earth’s insolation.

The dissociation rate, J (s^{-1}), due to radiation in the wavelength range (λ_a, λ_b) is then given by

$$J = \int_{\lambda_a}^{\lambda_b} F_p(\lambda)\sigma(\lambda)\sum_i q_i(\lambda)d\lambda. \quad (3)$$

Integrating from the molecule’s ionization threshold wavelength, λ_{ion} , to ∞ gives the total dissociation rate due to irradiation from stellar photons (neglecting any dissociations from photons with sufficient energy to ionize the molecule).

To examine the importance of various spectral regions and features to dissociating a given molecule, a useful quantity is the “cumulative distribution” of j , normalized by the total dissociation rate, i.e.,

$$\tilde{j} = \frac{\int_{\lambda_{\text{ion}}}^{\lambda} j(\lambda')d\lambda'}{\int_{\lambda_{\text{ion}}}^{\infty} j(\lambda')d\lambda'}. \quad (4)$$

where λ_{ion} is the ionization threshold of the molecule. Physically, \tilde{j} expresses the fraction of the total photodissociation rate of a molecule that is due to photons with wavelengths between the molecule’s ionization threshold and λ . This permits easy visual interpretation from plotted \tilde{j} values of the degree to which a portion of the stellar spectrum is driving photodissociation of a molecule. The change in \tilde{j} over a given wavelength range gives the fraction of all dissociations attributable to photons in that range. Thus, ranges where \tilde{j} rapidly rises by a large amount indicate that stellar radiation with wavelengths in that range contributes disproportionately to the overall rate of photodissociation.

3.3.2. Comparison of Unshielded Photodissociation Rates of H_2 , N_2 , O_2 , O_3 , H_2O , CO , CO_2 , CH_4 , and N_2O among the MUSCLES Stars

For all 11 MUSCLES hosts and the Sun, we have tabulated the total unshielded dissociation rate, J , of the molecules H_2 , N_2 , O_2 , O_3 , H_2O , CO , CO_2 , CH_4 , and N_2O in Table 5. We considered these molecules because of their ubiquity and biological relevance. Among the MUSCLES stars, unshielded photodissociation rates of most of these molecules range over roughly an order of magnitude. Exceptions are H_2 , for which J values vary by over two orders of magnitude, and H_2O and CH_4 , for which J values vary by about a factor of five. Confining the comparison to only the 4 K stars or only the 7 M stars reduces the spread in J values, but the values still vary by a factor of a few or more. This variability between stars emphasizes the importance of spectral observations of individual exoplanet host stars to photochemical modeling.

Median unshielded photodissociation rates for the MUSCLES stars are generally a factor of a few higher than the Sun. However, one key exception to this is O_3 , for which the Sun’s relatively high level of photospheric flux in the NUV makes it the strongest dissociator by a factor of a few to nearly two orders of magnitude. We reserve further discussion of this molecule for Section 3.3.3.

Table 5
Dissociation Rates (s^{-1}) for Unshielded Molecules Receiving Bolometric Flux Equivalent to Earth's Insolation

Star	H ₂ $\times 10^{-7}$	N ₂ $\times 10^{-6}$	O ₂ $\times 10^{-6}$	O ₃ $\times 10^{-4}$	H ₂ O $\times 10^{-5}$	CO $\times 10^{-6}$	CO ₂ $\times 10^{-5}$	CH ₄ $\times 10^{-5}$	N ₂ O $\times 10^{-7}$	O ₂ /O ₃ ^a
GJ 1214	0.36	3.3	0.86	1.2	3.2	2.6	0.7	4.2	0.82	0.0069
GJ 876	0.49	2.6	4.2	1.4	3.0	2.0	0.54	4.1	2.4	0.031
GJ 436	1.2	3.7	1.2	1.9	3.3	2.6	0.69	4.5	0.85	0.0065
GJ 581	0.18	1.4	0.59	1.5	1.6	1.1	0.3	2.2	0.34	0.0039
GJ 667C	3.1	8.9	1.6	3.0	7.6	6.2	1.6	10	1.1	0.0052
GJ 176	2.1	5.1	3.4	2.6	4.2	3.4	0.89	5.6	2.2	0.013
GJ 832	2.2	6.5	2.1	3.7	5.7	4.6	1.2	7.8	1.6	0.0057
HD 85512	6.6	8.1	1.1	5.6	3.9	4.5	1.0	5.4	1.1	0.002
HD 40307	17	17	1.1	15	5.5	8.0	1.6	7.5	1.4	0.00075
HD 97658	17	15	1.8	34	4.9	7.2	1.4	6.6	3.6	0.00052
ϵ Eri	40	33	5.0	29	8.4	15	2.7	11	5.4	0.0017
Sun	0.56	1.1	2.5	81	1.1	0.67	0.13	0.88	13	0.0003

Note.

^a Ratio of the O₃ to O₂ dissociation rates.

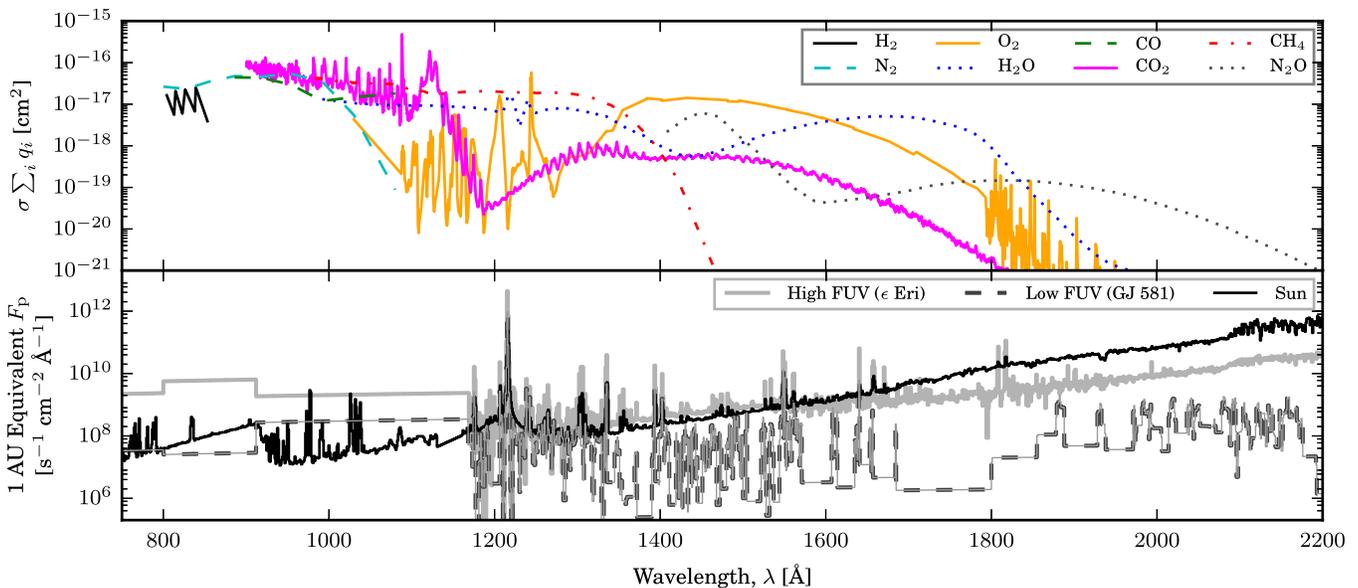


Figure 8. Top: photodissociation cross sections of the examined molecules. Bottom: SEDs of the reference stars: the most active MUSCLES star (as defined by the ratio of FUV to bolometric flux), least active MUSCLES star, and the Sun, converted to photon flux density instead of energy flux density and scaled such that the bolometric energy flux is equivalent to Earth's insolation.

To examine the wavelength dependence of unshielded photodissociation rates, we computed \tilde{j} curves for three representative stars: the most active star in the MUSCLES sample, ϵ Eri; the least active star, GJ 581; and the Sun, where we define activity as the ratio of the integrated FUV flux to the bolometric flux. Before presenting plots of the computed \tilde{j} curves, we first plot in Figure 8 the photodissociation cross sections incorporating quantum yields, $\sigma \sum_i q_i(\lambda)$, and the stellar photon flux density, $F_p(\lambda)$. The spectrum of GJ 581 had very low S/N in some areas, resulting in many spectral bins with negative flux density values due to the subtraction of background count rate estimates. Thus, for all stars we merged any negative-flux bins with their neighbors until the summed flux was above zero. This amounts to sacrificing resolution for S/N in these areas. The integrated and normalized product of the two curves in Figure 8 then gives the \tilde{j} curves plotted in Figure 9. Because \tilde{j} values are normalized by the overall photodissociation rate, J for a given input SED,

Figure 9 does not show how J varies between stars. However, this information is available in Table 5.

From Figure 9 it is clear that, with the exception of H₂O and N₂O, the same portions of the stellar spectra drive unshielded photodissociations of each molecule. This results from the relatively small differences in the reference star SED shapes over the range of photon wavelengths capable of photodissociating these molecules.

Although they are similar between stars for most molecules, the \tilde{j} curves reveal some intriguing structure. The dissociation of H₂, N₂, CO, and CO₂ is primarily due to flux shortward of the 1170 Å cutoff of the COS data. There is an apparent difference between the curves for the Sun versus GJ 581 and ϵ Eri in this region: the curve for the Sun shows jumps not present in the curves of GJ 581 and ϵ Eri. The jumps are caused by solar line emission that significantly dissociates these molecules. The jumps are not present in the curves for GJ 581 and ϵ Eri because these stars were not directly observed at

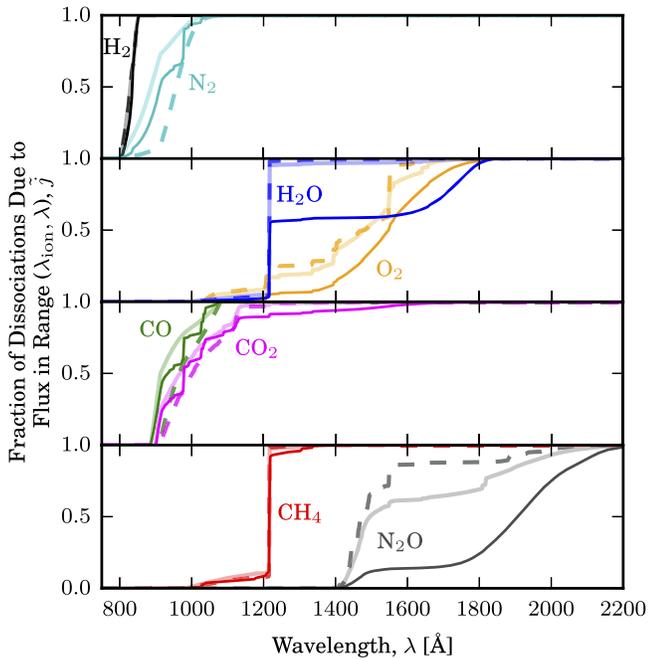


Figure 9. Cumulative photodissociation spectra. The curves show what fraction of the dissociation is due to photons with wavelengths from the molecule’s ionization threshold (λ_{ion}) to λ . This corresponds to the product of the curves in Figure 8 integrated from λ_{ion} to λ and normalized by the full integral. The curve’s rate of growth indicates the importance of that region of the spectrum to photodissociation of the molecule. Each molecule has three curves corresponding to the three reference stars whose spectra are plotted in the bottom panel of Figure 8 using the same line styles. Groups of curves are labeled by molecule and colored to match the top panel of Figure 8.

those wavelengths. Flux in this region is instead given by the broadband estimates described in Section 2.2 that lack any individual spectral lines. However, it is probable that line emission from GJ 581 and ϵ Eri in these ranges will, like the Sun, be important to the photodissociation of H_2 , N_2 , CO , and CO_2 . Similarly, note that while the CO_2 \tilde{j} curves of ϵ Eri and GJ 581 show some jumps, these are due to spikes in the photodissociation cross section of CO_2 rather than structure in the ϵ Eri and GJ 581 SEDs.

For O_2 , the photodissociation cross section has a narrow peak near $\text{Ly}\alpha$ and a broad, highly asymmetric peak at ~ 1400 Å. Thus, for the low-mass stars, emission from the $\text{Ly}\alpha$; $\text{Si IV } \lambda\lambda 1393, 1402$; and $\text{C IV } \lambda\lambda 1548, 1550$ lines contributes significantly to the overall photodissociation of O_2 . These are weaker relative to the FUV continuum in the Sun, so instead the solar FUV continuum dominates photodissociation.

Unshielded photodissociation of H_2O is dominated by $\text{Ly}\alpha$ in ϵ Eri and GJ 581 and unshielded photodissociation of CH_4 is dominated by $\text{Ly}\alpha$ for all three reference stars. For the Sun, dissociation of H_2O is roughly evenly shared between $\text{Ly}\alpha$ and the 1600–1800 Å FUV because of weaker $\text{Ly}\alpha$ emission and stronger FUV continuum emission. However, when interpreting the effect of $\text{Ly}\alpha$ emission on photodissociation rates, it is especially important to consider the assumption of unshielded molecules. $\text{Ly}\alpha$ is significantly attenuated by only small amounts of intervening H_1 , H_2 , H_2O , or CH_4 . These species scatter and absorb $\text{Ly}\alpha$ so readily that their presence limits $\text{Ly}\alpha$ dissociation to the uppermost reaches of an atmosphere. For reference, unity optical depth for the center of the $\text{Ly}\alpha$ line occurs at a column density of $3 \times 10^{17} \text{ cm}^{-2}$ of H_2 from

scattering and column densities of $7 \times 10^{16} \text{ cm}^{-2}$ of H_2O or $5 \times 10^{16} \text{ cm}^{-2}$ of CH_4 from absorption. These densities correspond to pressures of order 10^{-9} bar for planets with surface gravities close to Earth’s. For an investigation of the effect of $\text{Ly}\alpha$ on mini-Neptune atmospheres, see Miguel et al. (2015).

Like H_2O , N_2O dissociation is driven by differing wavelength ranges when comparing the SED of the Sun with that of ϵ Eri and GJ 581. While the photodissociation cross section of N_2O peaks in the FUV at around 1450 Å, a secondary peak that is several times broader and some two orders of magnitude lower occurs in the NUV near 1800 Å. For ϵ Eri and GJ 581, flux levels are slow to rise across the region encompassing these peaks, so radiation at the 1450 Å primary peak dominates dissociations. In contrast, the solar spectrum rises very rapidly over the same range and beyond. Consequently, solar radiation from ~ 1800 – 2100 Å dominates the dissociation of N_2O .

Overlying material will attenuate the stellar flux density at all wavelengths within an atmosphere, but the degree of attenuation varies by many orders of magnitude across the spectrum. This modifies both the overall level and the spectral content of the ambient radiation field with atmospheric depth. In turn, the shapes of the \tilde{j} curves in Figure 9 change and values of J diminish as atmospheric depth is increased. Many atmospheric models incorporate a treatment of radiative transfer in order to address the change in the ambient radiation field with atmospheric depth and accurately incorporate photochemistry (see, e.g., Segura et al. 2005; Hu et al. 2012; Grenfell 2014; Rugheimer et al. 2015). Although inapplicable within an atmosphere, the unshielded photodissociation rates we present allow a physical comparison of “top of the atmosphere” photochemical forcing as a function of stellar host and wavelength.

3.3.3. Abiotic O_2 and O_3 Production and the Significance of Visible Radiation in O_3 Photodissociation

The photodissociation of O_2 and O_3 is of special interest because of the potential use of these molecules as biomarkers (e.g., Lovelock 1965; Selsis et al. 2002; Tian et al. 2014; Harman et al. 2015). Significant abiotic production of these molecules is possible in CO_2 -rich atmospheres through CO_2 dissociation that liberates free O atoms to combine with O and O_2 , creating O_2 and O_3 (Selsis et al. 2002; Hu et al. 2012; Domagal-Goldman et al. 2014; Tian et al. 2014; Harman et al. 2015). Abiotic O_2 and O_3 generation is also possible in water dominated atmospheres from H_2O photolysis and subsequent H loss (Wordsworth & Pierrehumbert 2014). Further, O_2 and O_3 buildup by water photolysis is especially important early in the life of a low-mass star when XUV and UV fluxes are higher (Luger & Barnes 2015). Hu et al. (2012) found that the strongest factor controlling abiotic O_2 levels is the presence of reducing species (such as outgassed H_2 and CH_4), extending the previous work by Segura et al. (2007) to planetary scenarios with very low volcanic outgassing rates. In addition Tian et al. (2014) and Harman et al. (2015) found that the ratio of FUV to NUV flux positively correlates with abiotic O_2 and O_3 abundances. This ratio is tabulated for the MUSCLES stars in Paper I. Tian et al. (2014) and Harman et al. (2015) suggest that the dependence of abiotic O_2 abundance on the FUV/NUV ratio results from a variety indirect pathways according to the various atmospheric compositions considered.

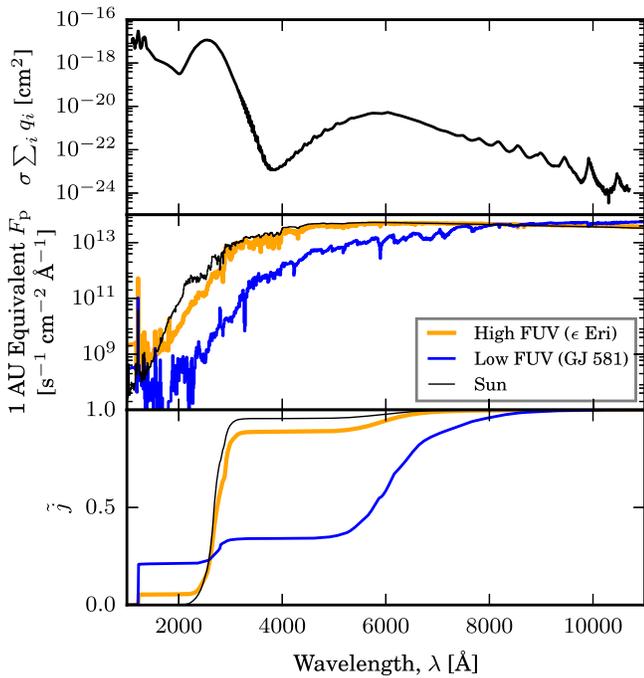


Figure 10. Same as Figures 8 and 9 for O_3 . Top: photodissociation cross section of O_3 . Middle: photon flux density spectra of ϵ Eri, GJ 581, and the Sun scaled such that the bolometric energy flux is equivalent to Earth’s insolation. Bottom: fraction of O_3 dissociations due to photons with wavelengths between λ_{ion} and λ if exposed to the spectra in the middle panel.

In contrast to the indirect effect the FUV/NUV ratio has on O_2 levels, it has a direct impact on the production of O_3 in CO_2 -rich or O_2 -rich atmospheres. NUV photons dissociate O_3 , whereas FUV photons liberate free O atoms from O_2 and CO_2 that can then combine in a three-body reaction with O_2 to create O_3 .

Photodissociation rates vary among MUSCLES stars by over an order of magnitude for O_2 and O_3 (Table 5). As a proxy for the relative forcing to higher O_3 populations in an O_2 -rich atmosphere, we give the ratio of the unshielded dissociation rate of O_2 to O_3 in Table 5. Higher ratios indicate photochemical forcing toward larger O_3 concentrations through more rapid dissociation of O_2 or less rapid dissociation of O_3 . These ratios vary by factors of a few in the MUSCLES stars, with ϵ Eri standing out as having a ratio roughly an order of magnitude below the median. The Sun has comparatively weak O_3 forcing—two orders of magnitude below the median.

Interestingly, NUV flux does not dominate the dissociation of O_3 in all stars. For those with weak levels of NUV flux, visible light can play a roughly equal role. This is apparent when examining photodissociation of O_3 by the unattenuated radiation of ϵ Eri and GJ 581. Figure 10 shows this, plotting the photodissociation cross section of O_3 , photon flux density of the reference stars, and the resulting \bar{j} curves for O_3 (akin to Figures 8 and 9). For ϵ Eri and the Sun, 2500–3200 Å NUV flux is responsible for the bulk of O_3 dissociation. In stark contrast, for GJ 581 about 20% of the dissociation is a result of Ly α , and visible photons are responsible for most of the remainder. The NUV flux of GJ 581, at roughly two orders of magnitude below that of the Sun and ϵ Eri, provides a minimal contribution, about 10%, to the photolysis of O_3 .

Indeed, visible radiation plays an important role in unshielded photodissociation of O_3 for all of the MUSCLES M dwarfs. This pattern is depicted in Figure 11, comparing the

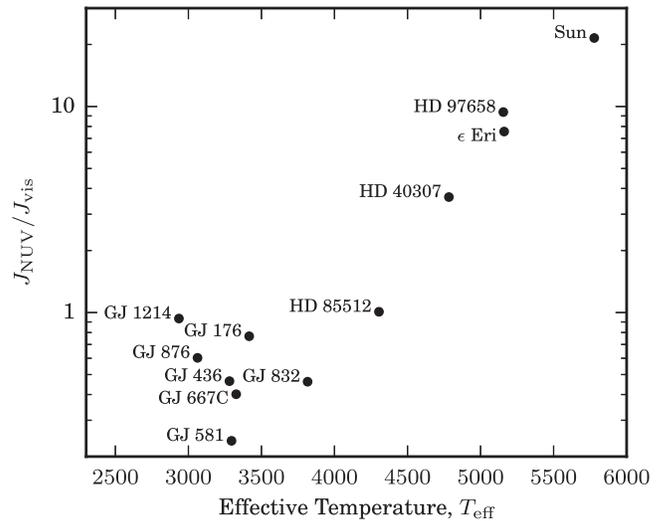


Figure 11. Ratio of dissociations of O_3 from NUV photons to those from visible photons if exposed to the unattenuated flux of the MUSCLES stars vs. the stellar effective temperatures we estimated. Visible flux is important for all stars with T_{eff} under 4500 K, but this importance diminishes with increasing T_{eff} above 4500 K.

ratio of O_3 photolysis from NUV flux to visible flux with the stellar effective temperature. For the M dwarfs, photospheric continuum flux is no longer a significant contributor in the NUV and visible radiation contributes as much or more than NUV flux to the photolysis of O_3 . This suggests that, among M dwarf hosts, the ratio of FUV to visible flux might prove to be a better predictor of relative O_3 populations in planetary atmospheres than the FUV/NUV ratio.

3.4. A Note on Models of Stellar Atmospheres

Section 3.3 demonstrates the importance of including accurate levels of the stellar FUV flux in the SEDs used as input to photochemical models. When such models involve a system lacking observations of the host star, the modeler must draw from other sources to provide the necessary input SED. One such source is the several existing codes that can synthesize a stellar SED by modeling the stellar atmosphere, most notably the PHOENIX code (Hauschildt & Baron 1999). However, most implementations of these codes truncate the stellar atmosphere above the photosphere. This includes the Husser et al. (2013) code that generated the spectra covering the visible and IR portions of the MUSCLES SEDs. Such photosphere-only models greatly underestimate levels of FUV flux emitted by low-mass, cool stars.

The extremity of this underestimate is illustrated in Figure 12 by plotting the panchromatic SED of GJ 832, incorporating the UV observations by *HST*, against the photosphere-only PHOENIX spectrum interpolated from the Husser et al. (2013) catalog. Comparing the two spectra shows that the Husser et al. (2013) PHOENIX spectrum predicts a flux lower than the observed value by three orders of magnitude or greater throughout the FUV. A similar discrepancy has been noted before for stars of mass equal to or less than that of the Sun (Franchini et al. 1998; Seager et al. 2013; Shkolnik & Barman 2014; Rugheimer et al. 2015). Because the FUV is where the photodissociation cross sections of many molecules peak, the use of inaccurately low FUV flux levels could underpredict rates of photodissociations if, like the MUSCLES stars, most

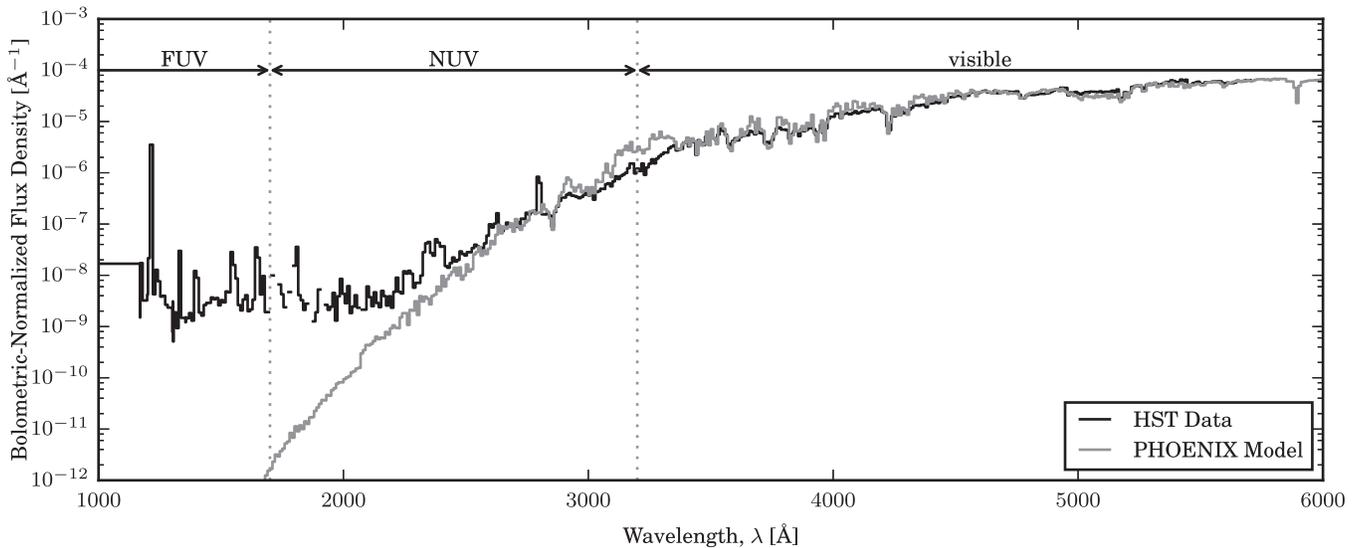


Figure 12. Fit of the PHOENIX model spectrum for GJ 832. The spectra are identically binned to facilitate comparison. The PHOENIX models closely track the visible spectrum measured by *HST*, but poorly predict the flux in the FUV and NUV. The discrepancy near 3200 Å dominates the integrated NUV flux, causing a factor ~ 2 overprediction of the NUV by the PHOENIX models. The PHOENIX FUV flux is negligible compared to that measured by *HST*.

low-mass stars show levels of UV emission exceeding that predicted by photosphere-only models. Similarly, photosphere-only models are likely to underpredict EUV levels and associated heating rates in planetary thermospheres by many orders of magnitude.

The prediction of lower than observed levels of short-wavelength flux is not an oversight of the Husser et al. (2013) PHOENIX model. Rather, the Husser et al. (2013) models and most others omit the stellar upper atmosphere by design. While not common, models that attempt to reproduce emission from the stellar upper atmosphere do exist. The same PHOENIX code often used for modeling photospheric emission has also been adapted for chromospheric emission from five M dwarfs (Fuhrmeister et al. 2005) and for CN Leonis during a flare (Fuhrmeister et al. 2010). In addition, much work on chromospheric modeling has been done by the Houdebine group, beginning with Houdebine & Doyle (1994). These works focused on ground-observable NUV and visible emission, rather than the FUV emission most important to photochemistry. Alternatively, a model by Fontenla et al. (2016) reproduces much of the spectral structure of the X-ray, UV, and visible emission of the M1.5 V star GJ 832, and work underway by Peacock et al. (2015) seeks to model the same emission for a variety of M dwarf stars. Generalized to a broader range of stellar types, we expect models like this will be indispensable for the simulation of photochemistry in the atmospheres of planets orbiting stars lacking detailed observations. These models require a sample of stars with detailed UV observations for validation. The MUSCLES Treasury Survey provides this sample.

4. SUMMARY

We have presented a catalog of SEDs spanning X-ray to IR wavelengths for 11 nearby, low-mass exoplanet host stars. Using these spectra, we examined the line versus continuum emission energy budget in the FUV, detecting an inter-line continuum at $>3\sigma$ confidence for 5 stars. This revealed a likely ionization edge structure in the FUV continuum of ϵ Eri that could represent the Si II recombination continuum. We found

that, when it was detected, the continuum contributed around 10% of the flux in the range spanned by the continuum bands. Uncertainties on nondetections were also consistent with this level.

Additionally, we examined photodissociation rates of the molecules H_2 , N_2 , O_2 , O_3 , H_2O , CO , CO_2 , CH_4 , and N_2O assuming that these species were exposed directly to the unattenuated stellar flux. We found total dissociation rates driven by different stars in our sample varied by over an order of magnitude for the majority of these molecules. By comparison, the photodissociation rates for the solar spectrum are near or below the lowest values for MUSCLES stars, except for O_3 where the solar photospheric NUV flux produces dissociation rates several times larger than the median value for MUSCLES stars. We also examined the relative importance of different spectral regions to the photodissociation of unshielded molecules driven by the SEDs of the most active host star (ϵ Eri), least active host star (GJ 581), and the Sun. There is little variation among these stars in which portion of the stellar emission is responsible for the bulk of unshielded photodissociations for most of the molecules examined. However, for O_3 we found that the dominant dissociative wavelength range can move from the NUV to the visible for low-mass stars with little NUV flux.

The spectral catalog presented in this paper, archived with MAST¹⁸, provides critical data for vetting stellar models and simulating photochemistry in planetary atmospheres.

The data presented here were obtained as part of the *HST* Guest Observing programs #12464 and #13650 as well as the COS Science Team Guaranteed Time programs #12034 and #12035. This work was supported by NASA grants HST-GO-12464.01 and HST-GO-13650.01 to the University of Colorado at Boulder. We thank Tom Woods and Chris Moore for useful discussions that provided a solar context to the work. This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the

¹⁸ <https://archive.stsci.edu/prepds/muscles/>

University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation. S.R. would like to acknowledge support from the Simons Foundation (339489, Rugheimer). F.T. is supported by the National Natural Science Foundation of China (41175039) and the Startup Fund of the Ministry of Education of China (20131029170). This work was partially supported by *Chandra* grants GO4-15014X and GO5-16155X from Smithsonian Astrophysical Observatory and NASA XMM grant NNX16AC09G.

REFERENCES

- Abrahamyan, H. V., Mickaelian, A. M., & Knyazyan, A. V. 2015, *A&C*, **10**, 99
- Ammons, S. M., Robinson, S. E., Strader, J., et al. 2006, *ApJ*, **638**, 1004
- Anderson, E., & Francis, C. 2012, *AstL*, **38**, 331
- Anglada-Escudé, G., & Butler, R. P. 2012, *ApJS*, **200**, 15
- Aparicio Villegas, T., Alfaro, E. J., Cabrera-Cañó, J., et al. 2010, *AJ*, **139**, 1242
- Arnold, K. A. 1996, in *ASP Conf. Ser.* 101, *Astronomical Data Analysis Software and Systems V*, ed. G. H. Jacoby, & J. Barnes (San Francisco, CA: ASP), 17
- Bailer-Jones, C. A. L. 2011, *MNRAS*, **411**, 435
- Barks, H. L., Buckley, R., Grieves, G. A., et al. 2010, *ChemBioChem*, **11**, 1240
- Bianchi, L., Herald, J., Efremova, B., et al. 2011, *Ap&SS*, **335**, 161
- Biretta, J., et al. 2015, *STIS Instrument Handbook, Version 14.0 (STScI)*
- Bonfanti, A., Ortolani, S., & Nascimbeni, V. 2016, *A&A*, **585**, A5
- Bonfils, X., Delfosse, X., Udry, S., et al. 2013, *A&A*, **549**, A109
- Bostroem, K., & Proffitt, C. 2011, *STIS Data Handbook, Version 6.0 (STScI)*
- Boyajian, T. S., von Braun, K., van Belle, G., et al. 2012, *ApJ*, **757**, 112
- Bryden, G., Beichman, C. A., Carpenter, J. M., et al. 2009, *ApJ*, **705**, 1226
- Carter, J. A., Winn, J. N., Holman, M. J., et al. 2011, *ApJ*, **730**, 82
- Cohen, O., Drake, J. J., Glöcker, A., et al. 2014, *ApJ*, **790**, 57
- Cowan, N. B., Greene, T., Angerhausen, D., et al. 2015, *PASP*, **127**, 311
- Cutri, R. M., Skrutskie, M. F., van Dyk, S., et al. 2003, *yCat*, 2246
- Cutri, R. M., et al. 2012, *yCat*, 2311
- Cutri, R. M., et al. 2014, *yCat*, 2328
- de Bruijne, J. H. J., & Eilers, A.-C. 2012, *A&A*, **546**, A61
- Delfosse, X., Bonfils, X., Forveille, T., et al. 2013, *A&A*, **553**, A8
- Deming, D., Seager, S., Winn, J., et al. 2009, *PASP*, **121**, 952
- Dere, K. P., Landi, E., Young, P. R., et al. 2009, *A&A*, **498**, 915
- Domagal-Goldman, S. D., Segura, A., Claire, M. W., Robinson, T. D., & Meadows, V. S. 2014, *ApJ*, **792**, 90
- Dressing, C. D., & Charbonneau, D. 2013, *ApJ*, **767**, 95
- Dressing, C. D., & Charbonneau, D. 2015, *ApJ*, **807**, 45
- Ducati, J. R. 2002, *yCat*, 2237
- Ehrenreich, D., Bourrier, V., Wheatley, P. J., et al. 2015, *Natur*, **522**, 459
- Eiroa, C., Marshall, J. P., Mora, A., et al. 2013, *A&A*, **555**, A11
- Fabricsius, C., Høg, E., Makarov, V. V., et al. 2002, *A&A*, **384**, 180
- Finch, C. T., Zacharias, N., Subasavage, J. P., Henry, T. J., & Riedel, A. R. 2014, *AJ*, **148**, 119
- Fontenla, J. M., Curdt, W., Haberleiter, M., Harder, J., & Tian, H. 2009, *ApJ*, **707**, 482
- Fontenla, J. M., Withrod, J., France, K., et al. 2016, *ApJ*, submitted
- Fossati, L., Ayres, T. R., Haswell, C. A., et al. 2013, *ApJL*, **766**, L20
- Fossati, L., Haswell, C. A., Froning, C. S., et al. 2010, *ApJL*, **714**, L222
- Fox, A. J., et al. 2015, *COS Data Handbook, Version 3.0 (STScI)*
- France, K., Froning, C. S., Linsky, J. L., et al. 2013, *ApJ*, **763**, 149
- France, K., Loyd, R. O. P., Youngblood, A., et al. 2016, *ApJ*, **820**, 89
- France, K., Schindhelm, E., Bergin, E. A., Roueff, E., & Abgrall, H. 2014, *ApJ*, **784**, 127
- Franchini, M., Morossi, C., & Malagnini, M. L. 1998, *ApJ*, **508**, 370
- Fressin, F., Torres, G., Charbonneau, D., et al. 2013, *ApJ*, **766**, 81
- Fuhrmeister, B., Schmitt, J. H. M. M., & Hauschildt, P. H. 2005, *A&A*, **439**, 1137
- Fuhrmeister, B., Schmitt, J. H. M. M., & Hauschildt, P. H. 2010, *A&A*, **511**, A83
- Gaidos, E., Mann, A. W., Lépine, S., et al. 2014, *MNRAS*, **443**, 2561
- Girard, T. M., van Altena, W. F., Zacharias, N., et al. 2011, *AJ*, **142**, 15
- Gould, B. G. 1879, *yCat*
- Grenfell, J. L., Gebauer, S., v. Paris, P., Godolt, M., & Rauer, H. 2014, *P&SS*, **98**, 66
- Haakonsen, C. B., & Rutledge, R. E. 2009, *ApJS*, **184**, 138
- Harman, C. E., Schwieterman, E. W., Schottelkotte, J. C., & Kasting, J. F. 2015, *ApJ*, **812**, 137
- Hauschildt, P. H., & Baron, E. 1999, *JCoAM*, **109**, 41
- Henry, T. J., Jao, W.-C., Subasavage, J. P., et al. 2006, *AJ*, **132**, 2360
- Høg, E., Fabricius, C., Makarov, V. V., et al. 2000, *A&A*, **355**, L27
- Holmberg, J., Nordström, B., & Andersen, J. 2009, *A&A*, **501**, 941
- Houdebine, E. R. 2010, *MNRAS*, **407**, 1657
- Houdebine, E. R., & Doyle, J. G. 1994, *A&A*, **289**, 169
- Howard, A. W., Marcy, G. W., Bryson, S. T., et al. 2012, *ApJS*, **201**, 15
- Hu, R., & Seager, S. 2014, *ApJ*, **784**, 63
- Hu, R., Seager, S., & Bains, W. 2012, *ApJ*, **761**, 166
- Husser, T.-O., Wende-von Berg, S., Dreizler, S., et al. 2013, *A&A*, **553**, A6
- Jenkins, J. S., Ramsey, L. W., Jones, H. R. A., et al. 2009, *ApJ*, **704**, 975
- Jones, D. O., & West, A. A. 2016, *ApJ*, **817**, 1
- Kaltenegger, L., Sasselov, D., & Rugheimer, S. 2013, *ApJL*, **775**, L47
- Kerwin, B. A., & Remmele, R. L. 2007, *Journal of Pharmaceutical Sciences*, **96**, 1468
- Kharchenko, N. V. 2001, *KFNT*, **17**, 409
- Knutson, H. A., Dragomir, D., Kreidberg, L., et al. 2014, *ApJ*, **794**, 155
- Knutson, H. A., Madhusudhan, N., Cowan, N. B., et al. 2011, *ApJ*, **735**, 27
- Koen, C., Kilkenny, D., van Wyk, F., & Marang, F. 2010, *MNRAS*, **403**, 1949
- Koppapparu, R. K. 2013, *ApJL*, **767**, L8
- Kreidberg, L., Bean, J. L., Désert, J.-M., et al. 2014, *Natur*, **505**, 69
- Kulow, J. R., France, K., Linsky, J., & Loyd, R. O. P. 2014, *ApJ*, **786**, 132
- Lalitha, S., Poppenhaeger, K., Singh, K. P., Czesla, S., & Schmitt, J. H. M. M. 2014, *ApJL*, **790**, L11
- Lammer, H., Lichtenegger, H. I. M., Kulikov, Y. N., et al. 2007, *AsBio*, **7**, 185
- Lammer, H., Selsis, F., Ribas, I., et al. 2003, *ApJL*, **598**, L121
- Landolt, A. U. 2009, *AJ*, **137**, 4186
- Lawler, S. M., Beichman, C. A., Bryden, G., et al. 2009, *ApJ*, **705**, 89
- Lecavelier Des Etangs, A., Ehrenreich, D., Vidal-Madjar, A., et al. 2010, *A&A*, **514**, A72
- Lépine, S., & Gaidos, E. 2011, *AJ*, **142**, 138
- Linsky, J. L., Bushinsky, R., Ayres, T., Fontenla, J., & France, K. 2012, *ApJ*, **745**, 25
- Linsky, J. L., Fontenla, J., & France, K. 2014, *ApJ*, **780**, 61
- Linsky, J. L., Yang, H., France, K., et al. 2010, *ApJ*, **717**, 1291
- Lovelock, J. E. 1965, *Natur*, **207**, 568
- Luger, R., & Barnes, R. 2015, *AsBio*, **15**, 119
- Mamajek, E. E., & Hillenbrand, L. A. 2008, *ApJ*, **687**, 1264
- Matsui, T., & Abe, Y. 1986, *Natur*, **319**, 303
- Matsunaga, T., Hieda, K., & Nikaido, O. 1991, *Photochemistry & Photobiology*, **54**, 403
- Mermilliod, J.-C. 1986, *Catalogue of Eggen's UVB Data*, 0
- Miguel, Y., & Kaltenegger, L. 2014, *ApJ*, **780**, 166
- Miguel, Y., Kaltenegger, L., Linsky, J. L., & Rugheimer, S. 2015, *MNRAS*, **446**, 345
- Moses, J. I. 2014, *RSPTA*, **372**, 20130073
- Murray-Clay, R. A., Chiang, E. I., & Murray, N. 2009, *ApJ*, **693**, 23
- Myers, J. R., Sande, C. B., Miller, A. C., Warren, W. H., Jr., & Tracewell, D. A. 2015, *yCat*, 5145
- Neves, V., Bonfils, X., Santos, N. C., et al. 2013, *A&A*, **551**, A36
- Neves, V., Bonfils, X., Santos, N. C., et al. 2014, *A&A*, **568**, A121
- Patel, B. H., Percivalle, C., Ritson, D. J., Duffy, C. D., & Sutherland, J. D. 2015, *NatCh*, **7**, 301
- Peacock, S., Barman, T. S., & Shkolnik, E. 2015, in *American Astronomical Society Meeting Abstracts*, **225**, 138.26
- Pickles, A., & Depagne, É. 2010, *PASP*, **122**, 1437
- Reid, I. N., Cruz, K. L., Allen, P., et al. 2004, *AJ*, **128**, 463
- Ricker, G. R., Winn, J. N., Vanderspek, R., et al. 2014, *Proc. SPIE*, **9143**, 914320
- Ritson, D., & Sutherland, J. D. 2012, *NatCh*, **4**, 895
- Roeser, S., Demleitner, M., & Schilbach, E. 2010, *AJ*, **139**, 2440
- Rojas-Ayala, B., Covey, K. R., Muirhead, P. S., & Lloyd, J. P. 2012, *ApJ*, **748**, 93
- Röser, S., Schilbach, E., Schwan, H., et al. 2008, *A&A*, **488**, 401
- Rugheimer, S., Kaltenegger, L., Segura, A., Linsky, J., & Mohanty, S. 2015, *ApJ*, **809**, 57
- Salim, S., & Gould, A. 2003, *ApJ*, **582**, 1011
- Santos, N. C., Sousa, S. G., Mortier, A., et al. 2013, *A&A*, **556**, A150
- Seager, S., Bains, W., & Hu, R. 2013, *ApJ*, **777**, 95
- Seager, S., & Deming, D. 2010, *ARA&A*, **48**, 631
- Segura, A., Kasting, J. F., Meadows, V., et al. 2005, *AsBio*, **5**, 706

- Segura, A., Meadows, V. S., Kasting, J. F., Crisp, D., & Cohen, M. 2007, *A&A*, **472**, 665
- Selsis, F., Despois, D., & Parisot, J.-P. 2002, *A&A*, **388**, 985
- Senanayake, S. D., & Idriss, H. 2006, *PNAS*, **103**, 1194
- Shkolnik, E. L., & Barman, T. S. 2014, *AJ*, **148**, 64
- Shkolnik, E. L., Liu, M. C., Reid, I. N., Dupuy, T., & Weinberger, A. J. 2011, *ApJ*, **727**, 6
- Smith, R. K., Brickhouse, N. S., Liedahl, D. A., & Raymond, J. C. 2001, *ApJL*, **556**, L91
- Soubiran, C. 2010, *A&A*, **515**, A111
- Soubiran, C., Le Campion, J.-F., Cayrel de Strobel, G., & Caillo, A. 2010, *A&A*, **515**, A111
- Tevini, M. 1993, in *UV-B Radiation and Ozone Depletion: Effects on Humans, Animals, Plants, Microorganisms, and Materials*, ed. M. Tevini (Boca Raton, Florida: Lewis Publishers), 95
- The DENIS Consortium, 2005, *yCat*
- Tian, F., France, K., Linsky, J. L., Mauas, P. J. D., & Vieytes, M. C. 2014, *E&PSL*, **385**, 22
- Tian, F., Toon, O. B., Pavlov, A. A., & De Sterck, H. 2005, *ApJ*, **621**, 1049
- Torres, G., Winn, J. N., & Holman, M. J. 2008, *ApJ*, **677**, 1324
- TriAUD, A. H. M. J., Lanotte, A. A., Smalley, B., & Gillon, M. 2014, *MNRAS*, **444**, 711
- Tsantaki, M., Sousa, S. G., Adibekyan, V. Z., et al. 2013, *A&A*, **555**, A150
- Valenti, J. A., & Fischer, D. A. 2005, *ApJS*, **159**, 141
- Van Grootel, V., Gillon, M., Valencia, D., et al. 2014, *ApJ*, **786**, 2
- van Leeuwen, F. 2007, *A&A*, **474**, 653
- Vidal-Madjar, A., Lecavelier des Etangs, A., Désert, J.-M., et al. 2003, *Natur*, **422**, 143
- Voet, D., Gratzer, W. B., Cox, R. A., & Doty, P. 1963, *Biopolymers*, **1**, 193
- von Braun, K., Boyajian, T. S., Kane, S. R., et al. 2012, *ApJ*, **753**, 171
- von Braun, K., Boyajian, T. S., van Belle, G. T., et al. 2014, *MNRAS*, **438**, 2413
- Weis, E. W. 1996, *AJ*, **112**, 2300
- Winters, J. G., Henry, T. J., Lurie, J. C., et al. 2015, *AJ*, **149**, 5
- Woods, T. N., Chamberlin, P. C., Harder, J. W., et al. 2009, *GeoRL*, **36**, L01101
- Wordsworth, R., & Pierrehumbert, R. 2014, *ApJL*, **785**, L20
- Wright, J. T., Fakhouri, O., Marcy, G. W., et al. 2011, *PASP*, **123**, 412
- Yelle, R. V. 2004, *Icar*, **170**, 167
- Yoshino, K., Esmond, J. R., Cheung, A. S.-C., Freeman, D. E., & Parkinson, W. H. 1992, *P&SS*, **40**, 185
- Youngblood, A., France, K., Loyd, R. O. P., et al. 2016, *ApJ*, **824**, 101
- Zacharias, N., Finch, C. T., Girard, T. M., et al. 2012, *yCat*, 1322
- Zacharias, N., Monet, D. G., Levine, S. E., et al. 2004a, in *Bulletin of the American Astronomical Society, American Astronomical Society Meeting Abstracts*, **36**, 1418
- Zacharias, N., Urban, S. E., Zacharias, M. I., et al. 2004b, *AJ*, **127**, 3043